

# Cosmic Microwave Background Temperature Anisotropy from the Kinetic Sunyaev-Zel'dovich Effect

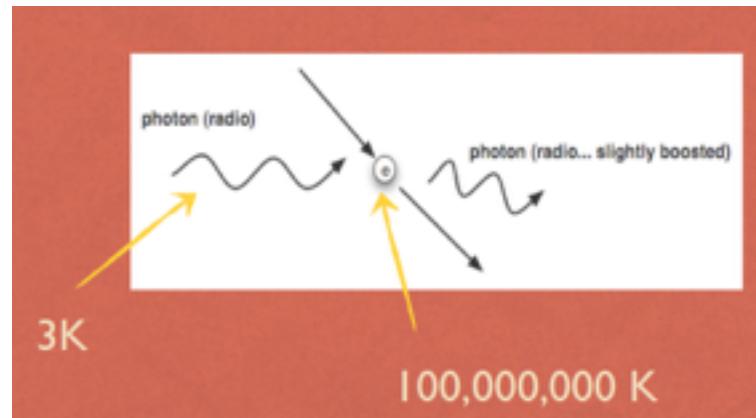
Speaker: Hyunbae Park (Technical Research Staff @ KASI)

## Collaborators:

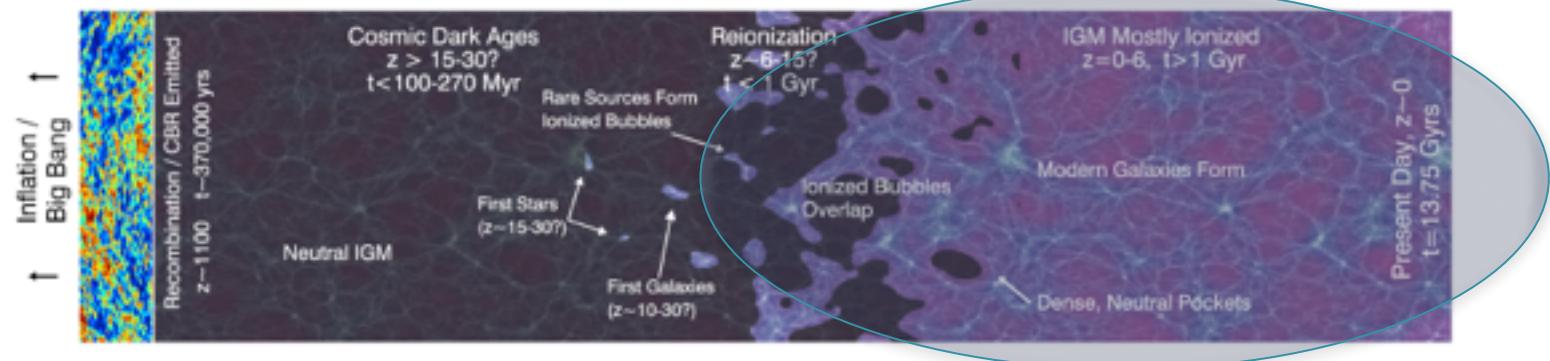
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Yi Mao (IAP), Jun Koda (Brera Astronomical Observatory),  
Marcelo Alvarez (CITA), J. Richard Bond (CITA),  
Kyungjin Ahn (Chosun Univ.), Ilian Iliev (U of Sussex),  
Garrelt Mellema (U of Stockholm)

# Intro) Sunyaev-Zel'dovich (SZ) Effect

Change of temperature due to interaction between CMB photons and free electrons...



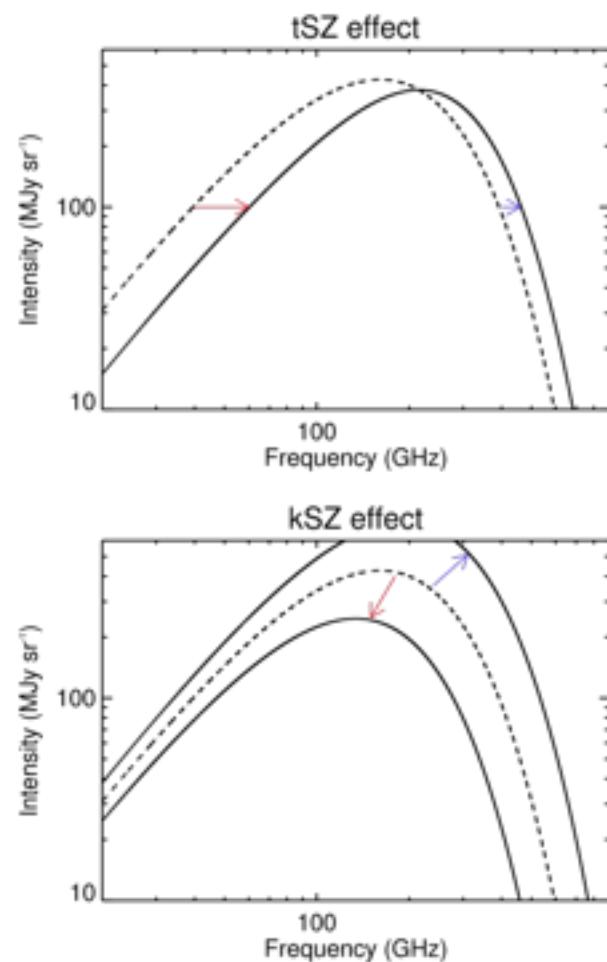
... produced by astrophysical sources in the post-Recombination era.



(Robertson et al. 2010)

# Intro) SZ Effect

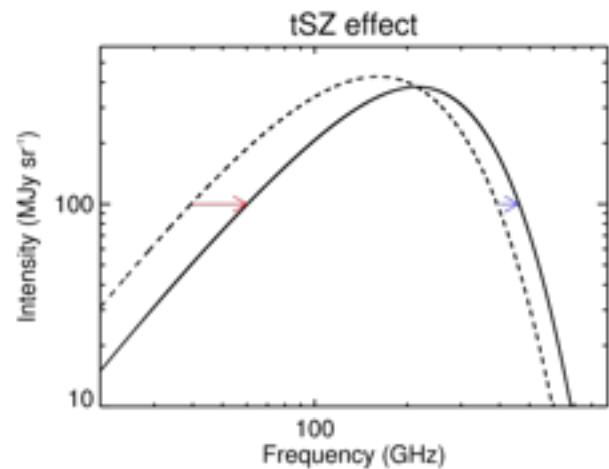
- Thermal SZ Effect (tSZ)  
Low energy photons up-scattered to have a higher frequency.
- Kinetic SZ Effect (kSZ)  
Doppler-shift by the bulk motion of the gas.



# Intro) SZ Effect

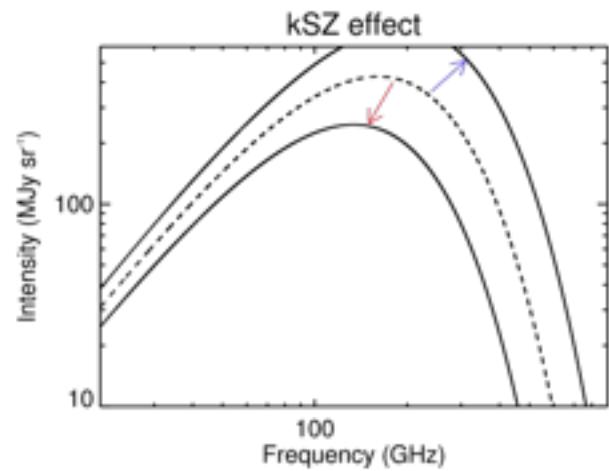
- Thermal SZ Effect (tSZ)

$$\frac{\Delta T}{T}(\hat{\gamma}) \propto \int d\tau \frac{kT}{m_e c^2}$$

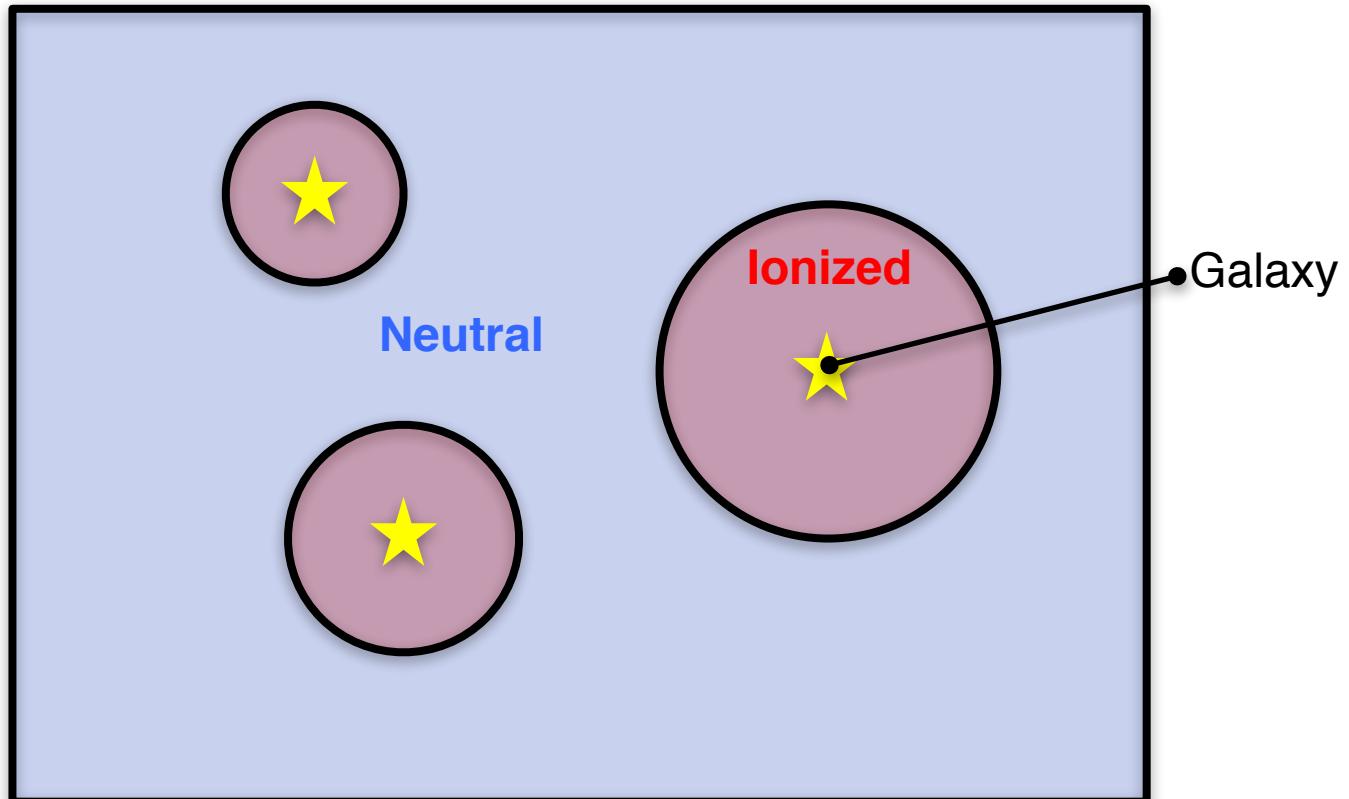


- Kinetic SZ Effect (kSZ)

$$\frac{\Delta T}{T}(\hat{\gamma}) = \int d\tau \hat{\gamma} \cdot \frac{\mathbf{v}}{c}$$

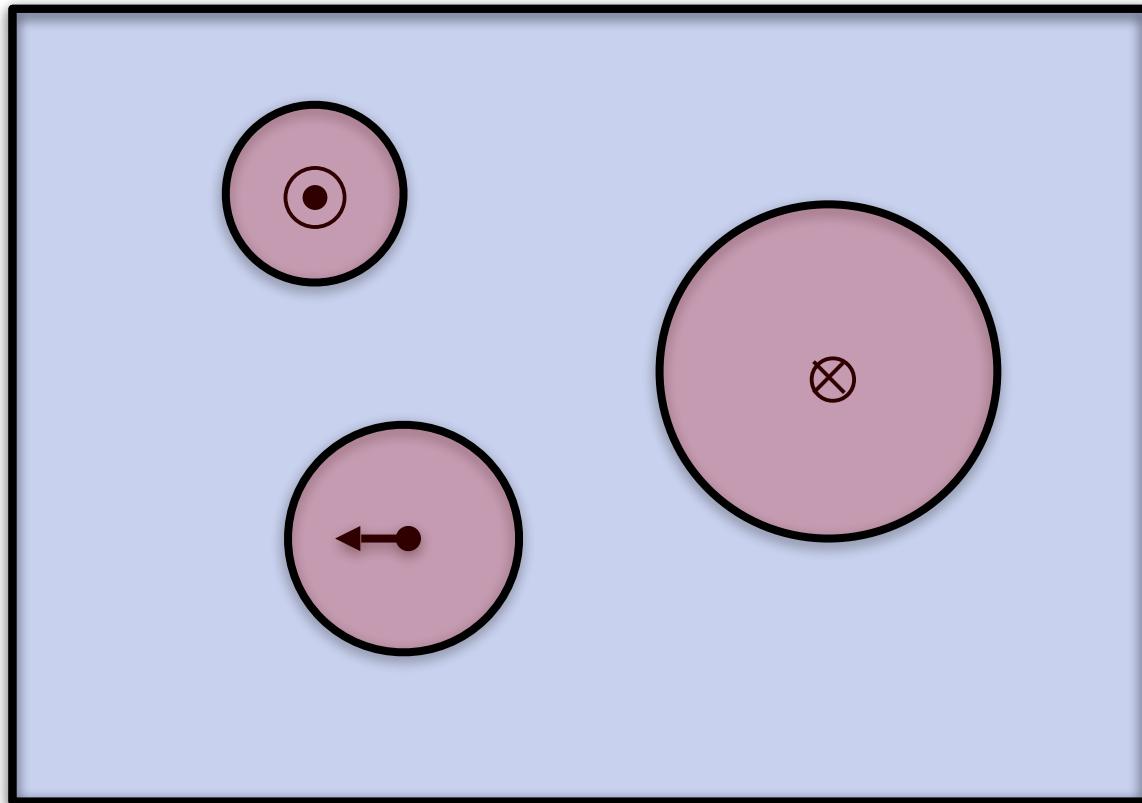


# Ionized IGM around Galaxies



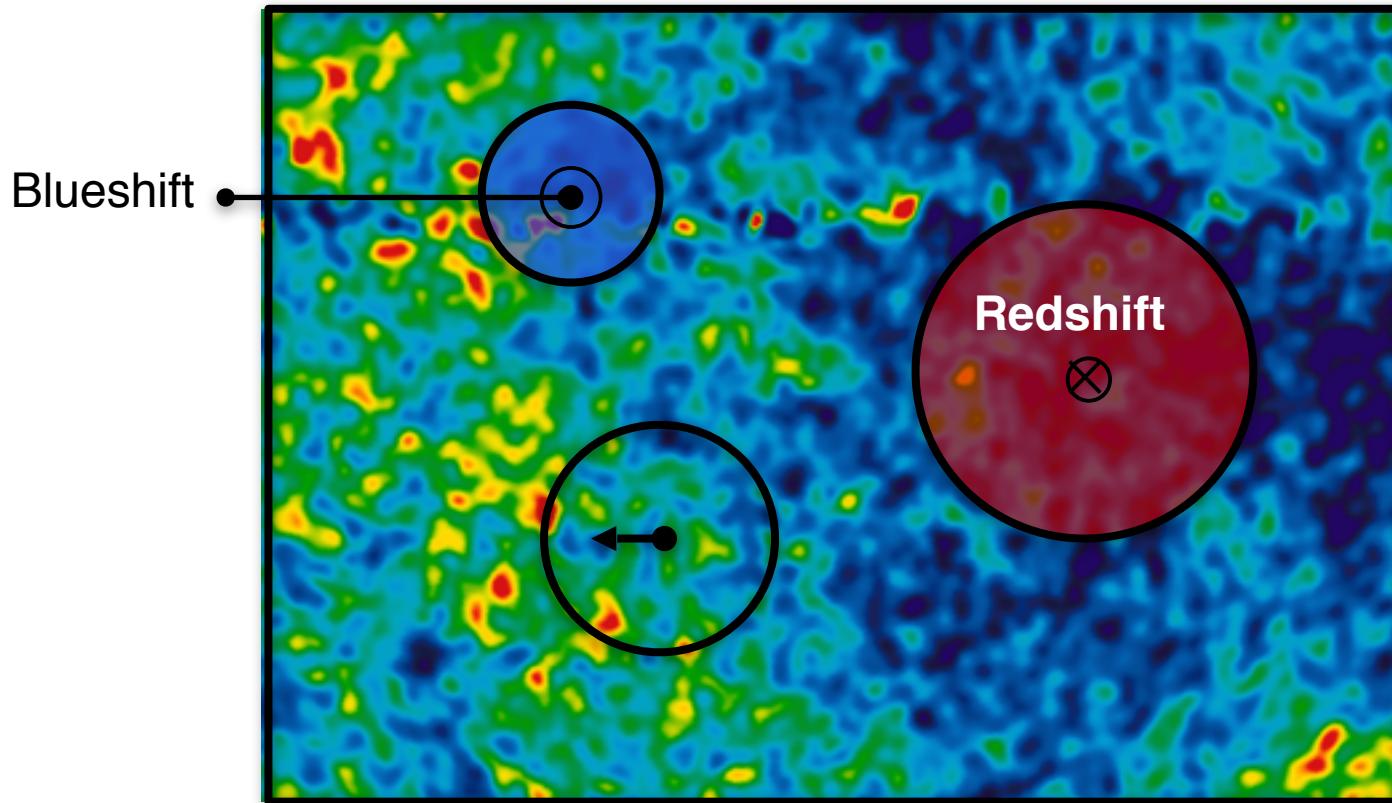
When the IGM has ionized parts, ...

# Ionized IGM with Peculiar Velocity



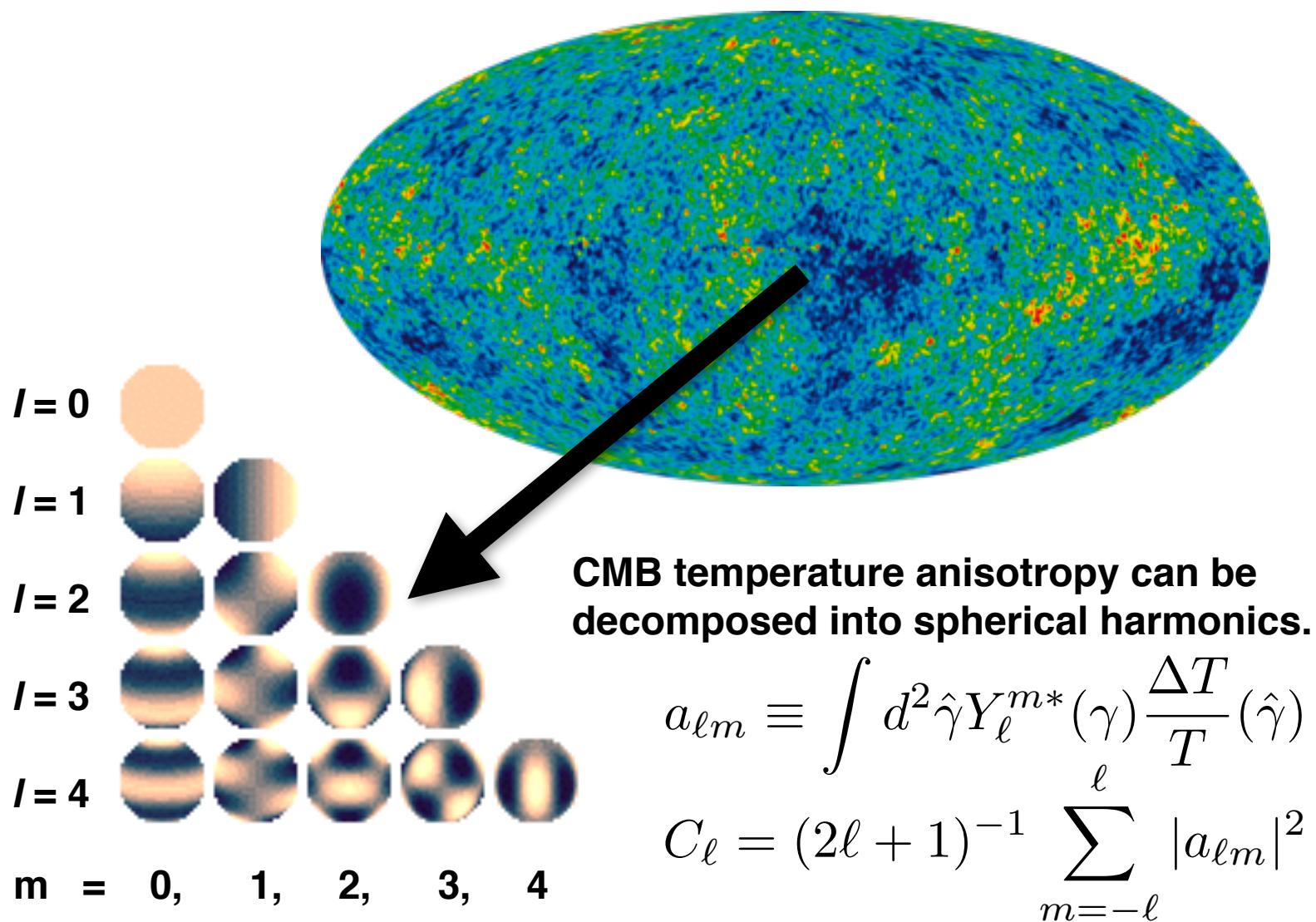
... each ionized patch can have different line-of-sight peculiar velocity.

# Kinetic SZ Effect

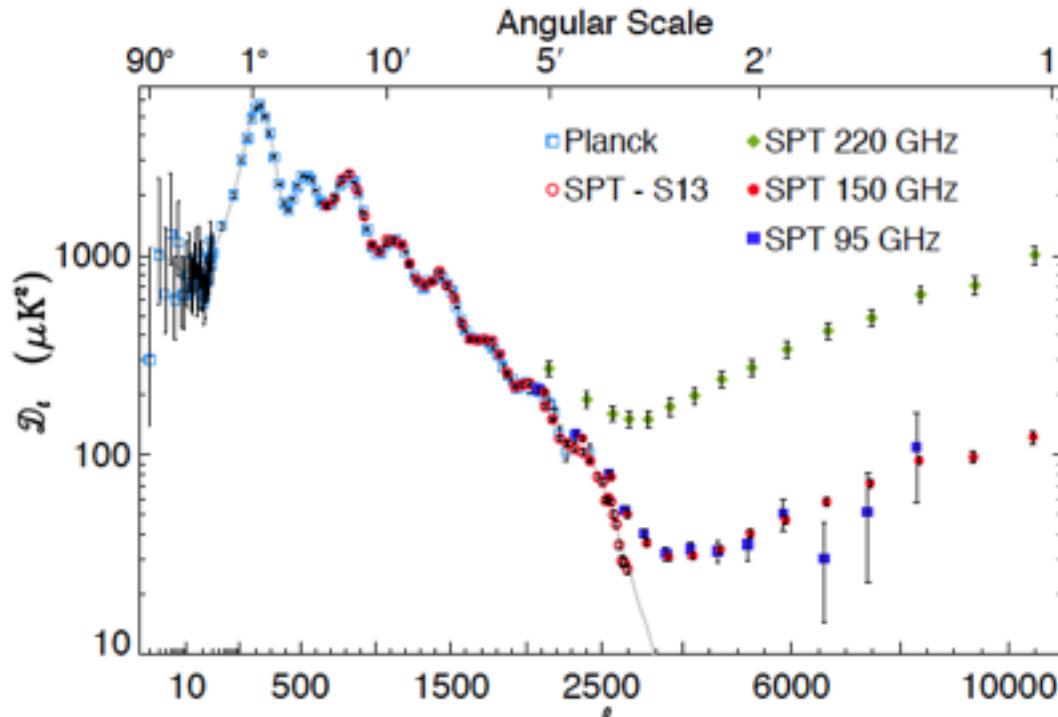


The kSZ effect would red/blueshift the CMB according to the line-of-sight velocity of the ionized IGM patches

# CMB Power Spectrum



# CMB Power Spectrum from the South-Pole Telescope

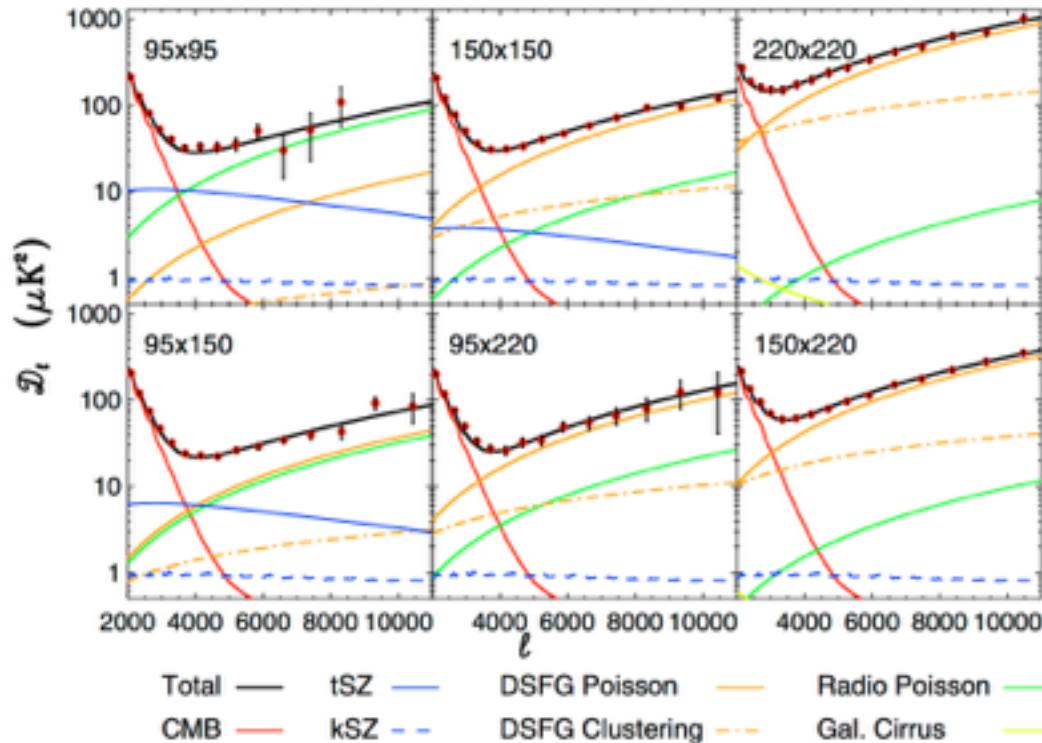


$$D_\ell \equiv (2\pi)^{-1} \ell(\ell+1) C_\ell$$

(George et al. 2015)

At  $\ell > 2000$  or at angular scales below 5 arcmin,  
contribution from astrophysical sources dominate the signal.

# kSZ Component in the CMB Power Spectrum



$$D_{\ell=3000}^{\text{kSZ}} = 2.9 \pm 1.3 \mu K^2 \quad (\text{George et al. 2015})$$

Using multi-frequency data, signals can be decomposed.  
For the kSZ effect, we only have rather a loose constraint  
at ( $\ell = 3000$ , 3 arcmin) at this point.

# kSZ power spectrum

$$\frac{\Delta T}{T}(\hat{\gamma}) = \int d\tau \hat{\gamma} \cdot \frac{\mathbf{v}}{c}$$

$$\mathbf{q} \equiv X(1 + \delta)\mathbf{v}$$

X: Ionized fraction

$$\frac{\Delta T}{T}(\hat{\gamma}) = \left( \frac{\bar{n}_{H,0}\sigma_T}{c} \right) \int \frac{ds}{a^2} (\mathbf{q} \cdot \hat{\gamma})$$

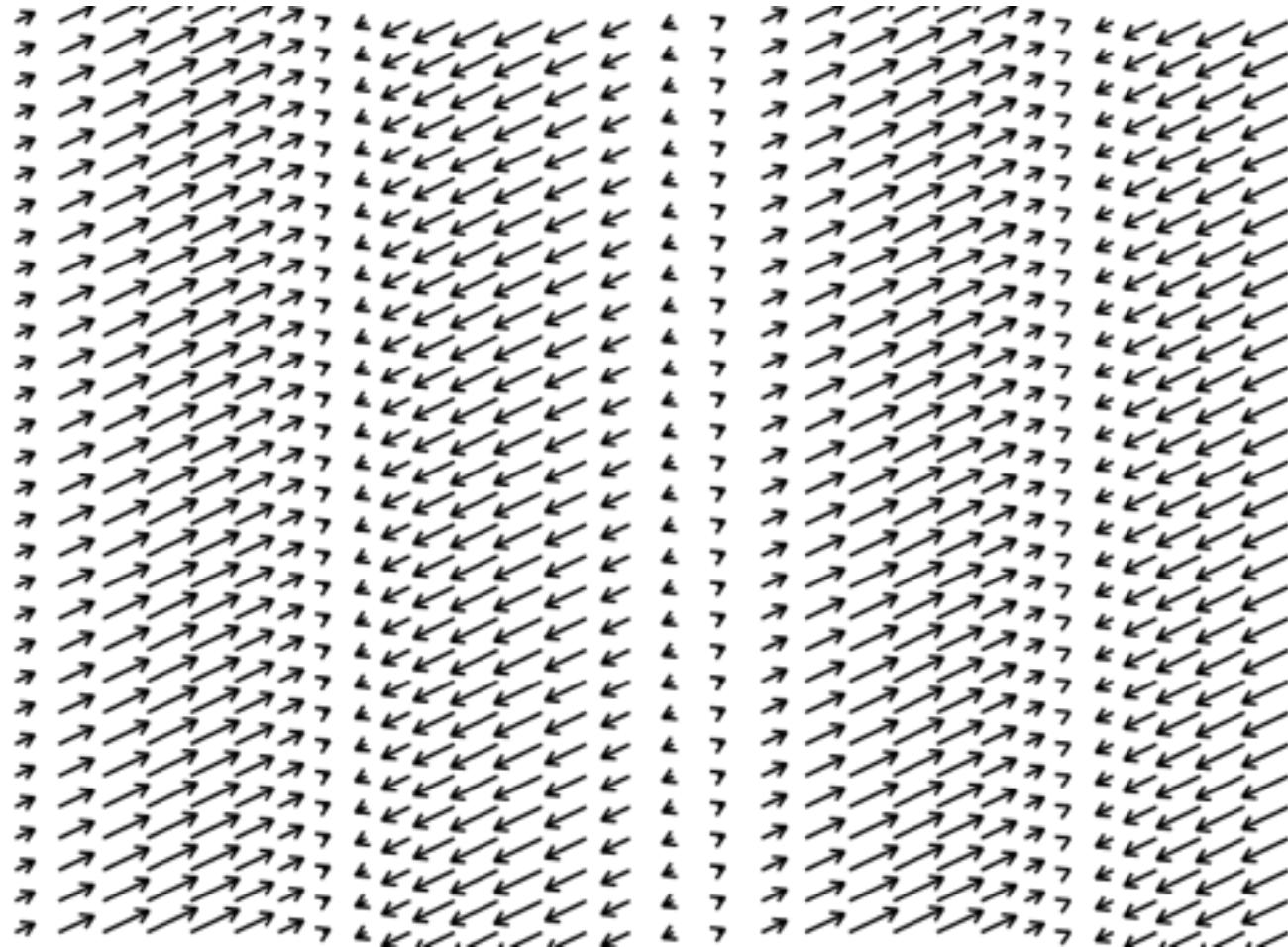
$$C_l = \left( \frac{\bar{n}_{H,0}\sigma_T}{c} \right)^2 \int \frac{ds}{s^2 a^4} \frac{P_{q\perp}(k = l/s, z)}{2}$$

(See Park et al. 2013 for the derivation)

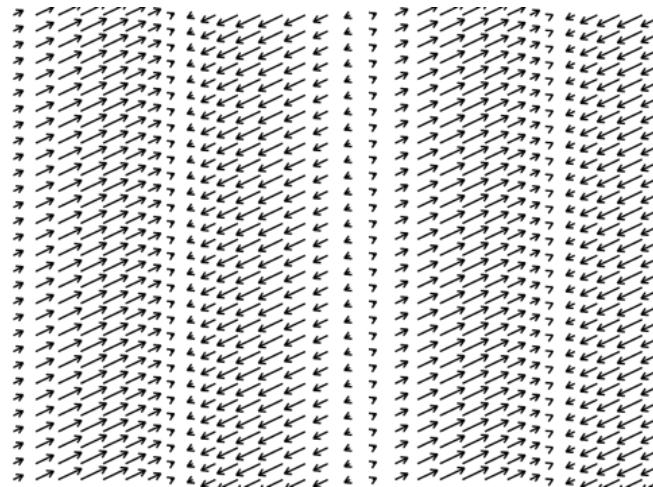
**kSZ power spectrum is given by the transverse component of ionized momentum field.**

# What is the “Transverse” mode?

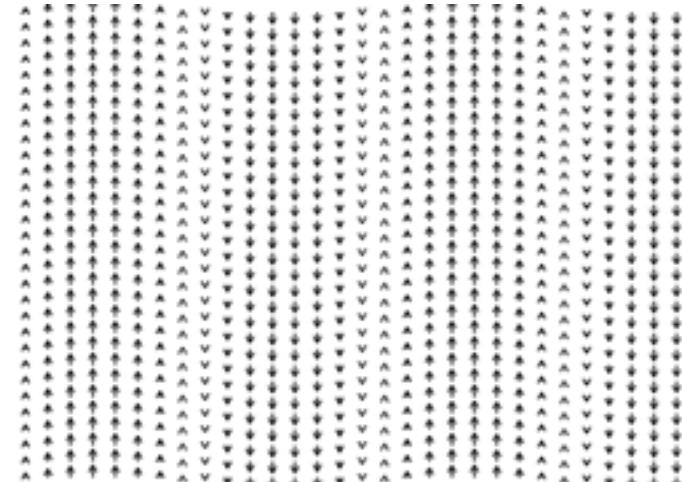
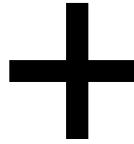
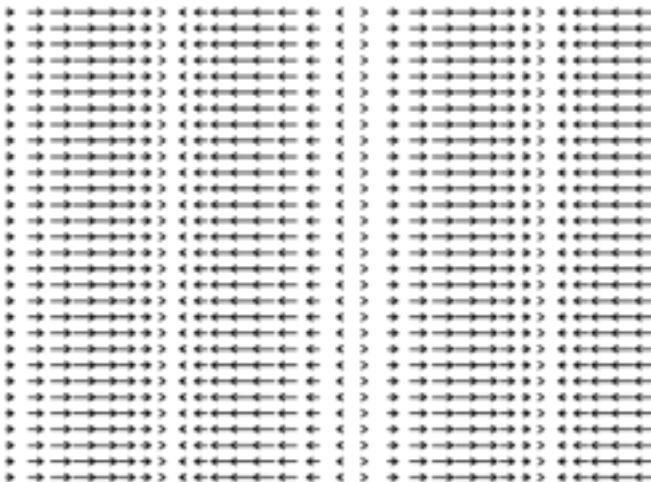
Here is a plane wave of a vector field.



The plane wave can be decomposed into...



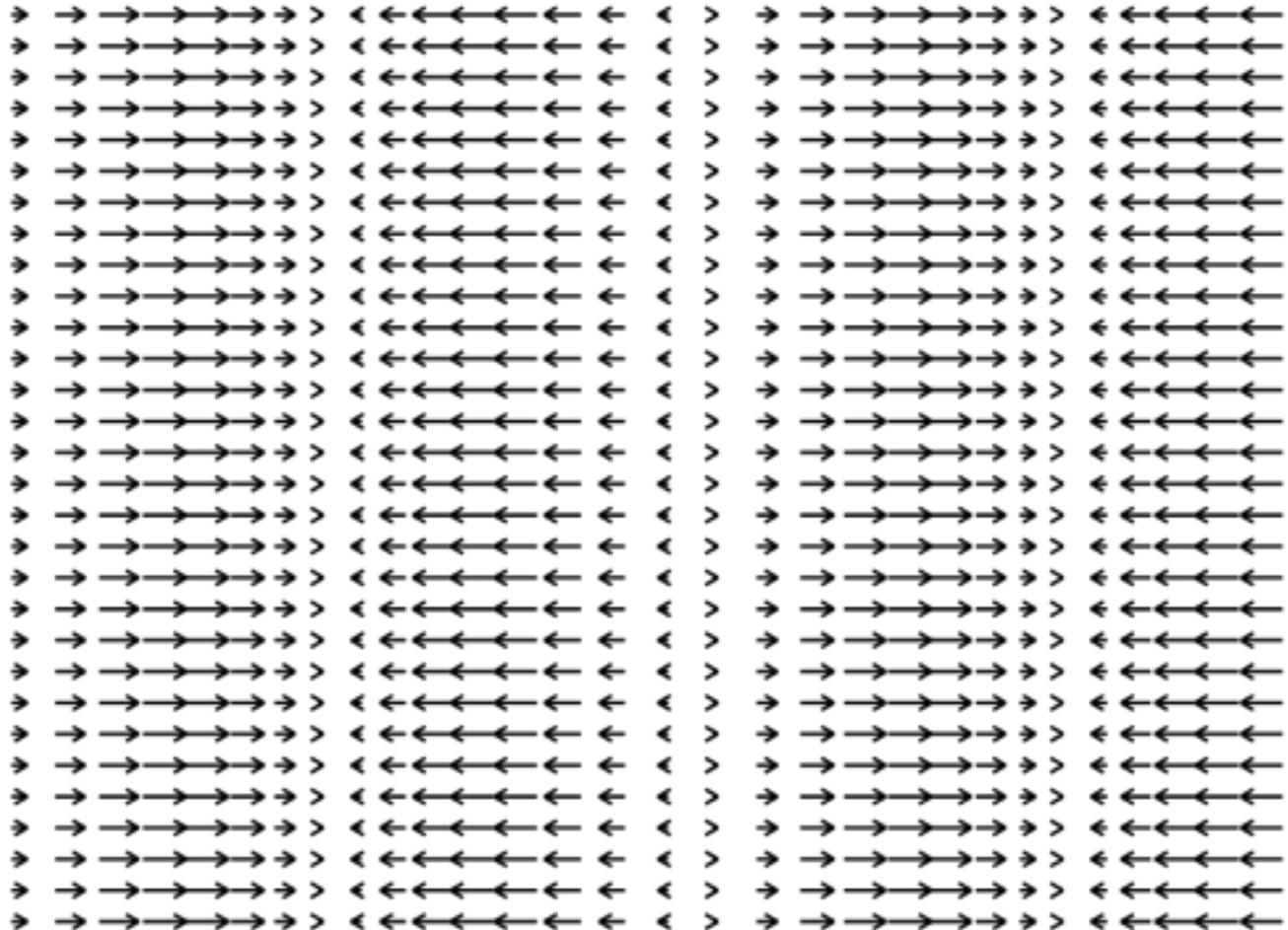
... the longitudinal mode and the transverse mode.



$$\tilde{\mathbf{q}}_{\parallel} \equiv \hat{\mathbf{k}} \cdot \tilde{\mathbf{q}}$$

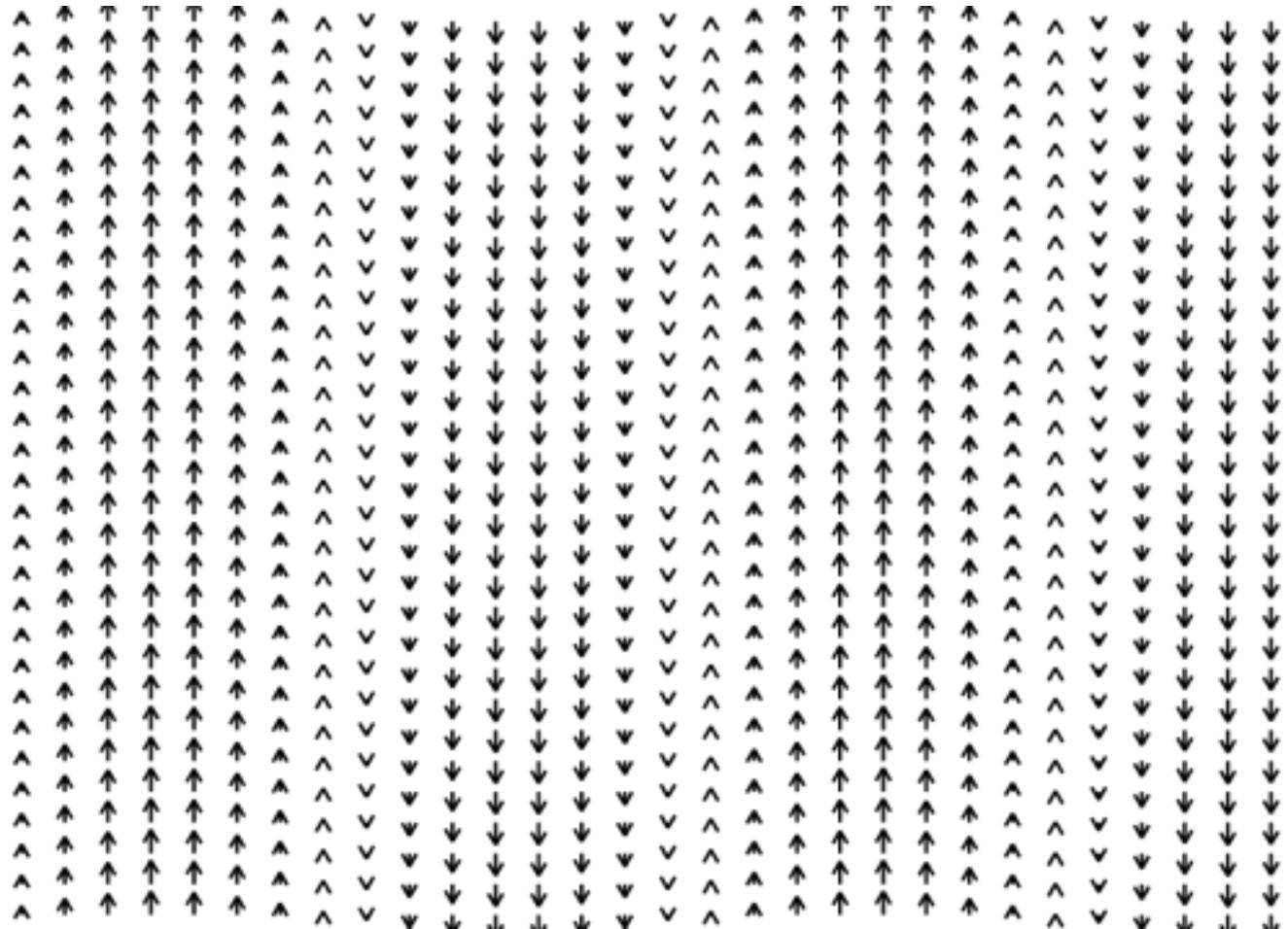
$$\tilde{\mathbf{q}}_{\perp} \equiv \tilde{\mathbf{q}} - \hat{\mathbf{k}} \cdot \tilde{\mathbf{q}}$$

# Longitudinal Mode



No kSZ contribution from the longitudinal mode!

# Transverse Mode



kSZ contribution when looking 

# kSZ power spectrum

$$\frac{\Delta T}{T}(\hat{\gamma}) = \int d\tau \hat{\gamma} \cdot \frac{\mathbf{v}}{c}$$

$$\mathbf{q} \equiv X(1 + \delta)\mathbf{v}$$

X: Ionized fraction

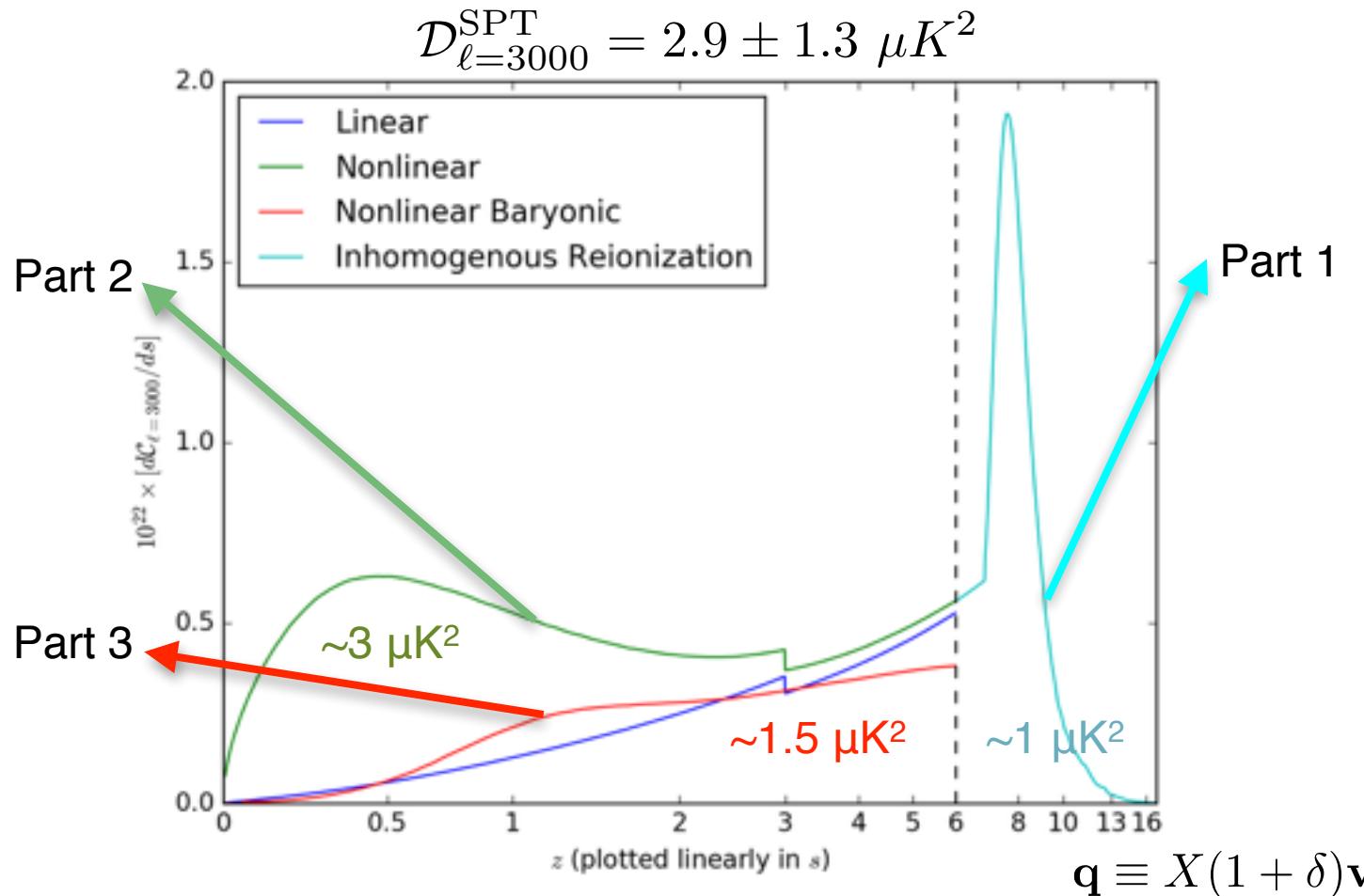
$$\frac{\Delta T}{T}(\hat{\gamma}) = \left( \frac{\bar{n}_{H,0}\sigma_T}{c} \right) \int \frac{ds}{a^2} (\mathbf{q} \cdot \hat{\gamma})$$

$$C_l = \left( \frac{\bar{n}_{H,0}\sigma_T}{c} \right)^2 \int \frac{ds}{s^2 a^4} \frac{P_{q\perp}(k = l/s, z)}{2}$$

(See Park et al. 2013 for the derivation)

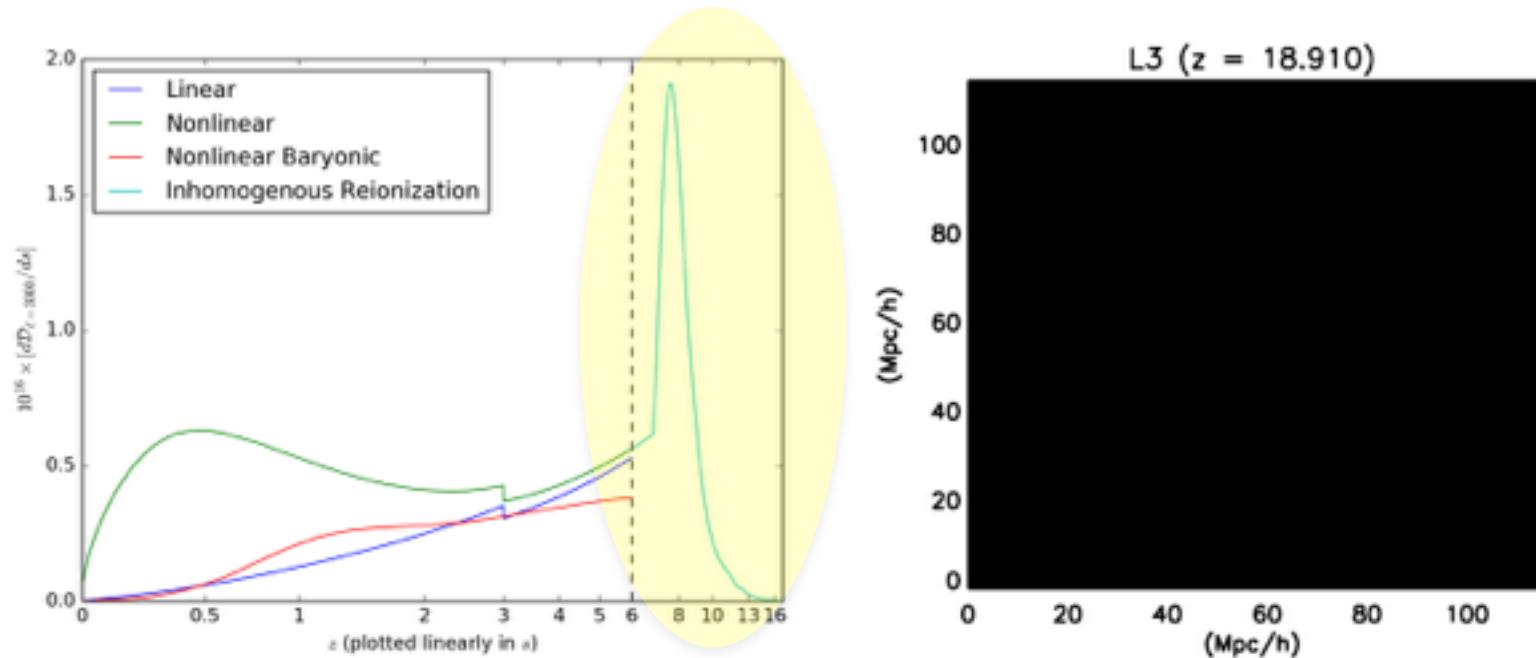
**kSZ power spectrum is given by projecting the transverse component of ionized momentum field along the line-of-sight.**

# Reionization vs Post-reionization



- 1) At  $z > 6$ , **inhomogeneously ionized IGM** gives a huge boost to the signal.
- 2) At  $z < 6$ , **non-linear growth of structure** and **baryonic physics** make huge impacts on the signal.

# Part 1) Reionization Signal



$$\mathbf{q} = X(1 + \delta)\mathbf{v} = X(1 + \delta_{lin})\mathbf{v}_{lin}$$

- 1) Inhomogeneous ionization strongly enhances the signal.
- 2) The signal is highly dependent on the properties of the source.

# Part 1) Properties of Ionizing Sources During EoR

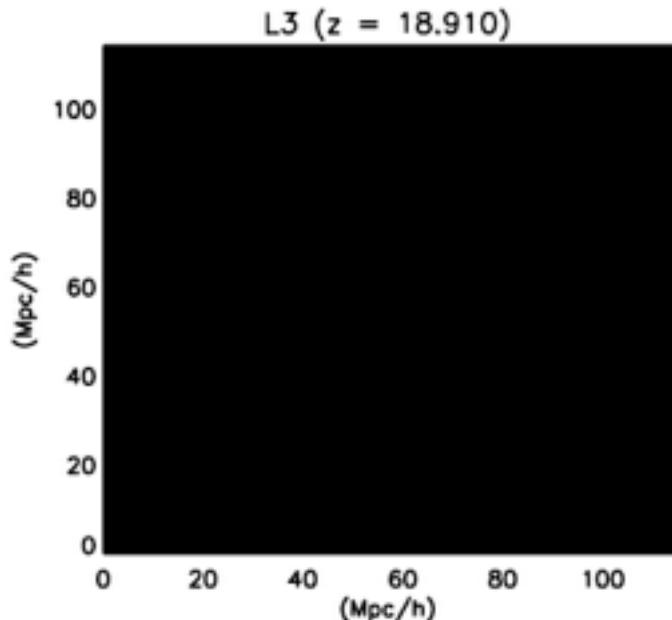
- High-mass Atomic Cooling Halos:  $M > 10^9 M_\odot$   
Never suppressed.
- Low-mass Atomic Cooling Halos:  $10^8 M_\odot < M < 10^9 M_\odot$   
Suppressed when surrounding region is ionized.
- Minihalos:  $M < 10^8 M_\odot$   
Suppressed when surrounding region is ionized or exposed to H<sub>2</sub> dissociating radiation.

# 1-1) Models with and without Low-mass Galaxies

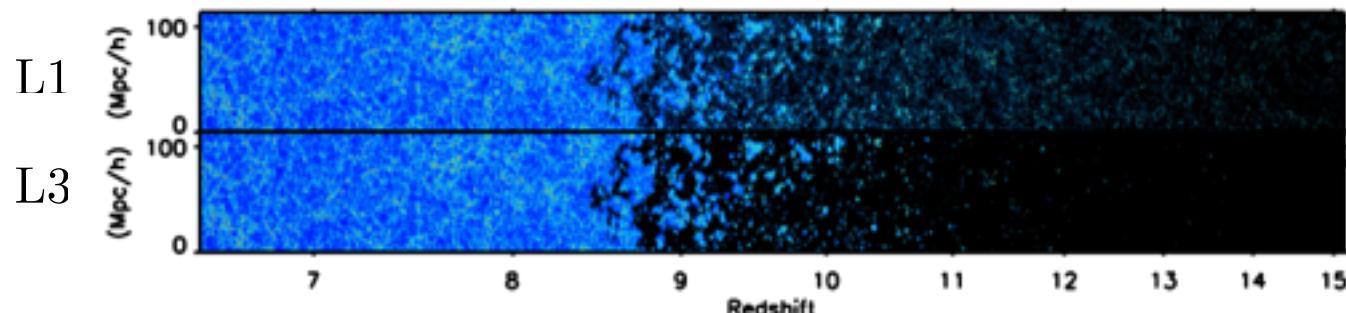
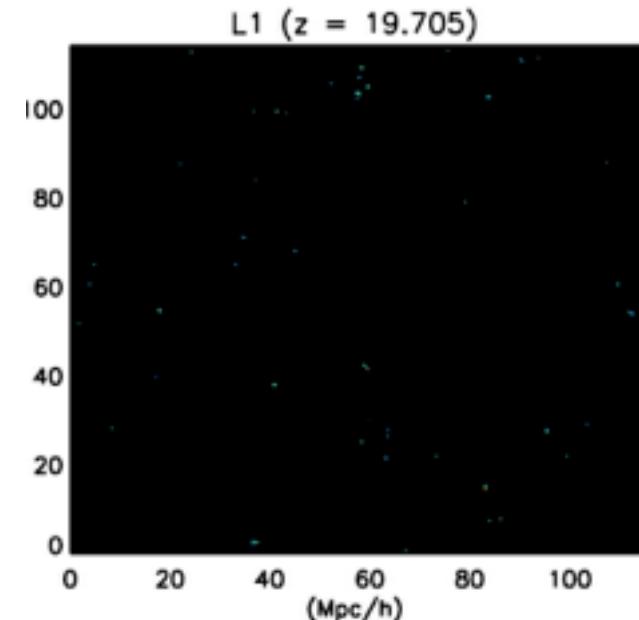
- **High-mass Atomic Cooling Halos:**  $M > 10^9 M_\odot$   
**Never suppressed.**
- **Low-mass Atomic Cooling Halos:**  $10^8 M_\odot < M < 10^9 M_\odot$   
**Suppressed when surrounding region is ionized.**
- **Minihalos:**  $M < 10^8 M_\odot$   
Suppressed when surrounding region is ionized or exposed to H<sub>2</sub> dissociating radiation.

# 1-1) Models with and without Low-mass Galaxies

High-mass sources only

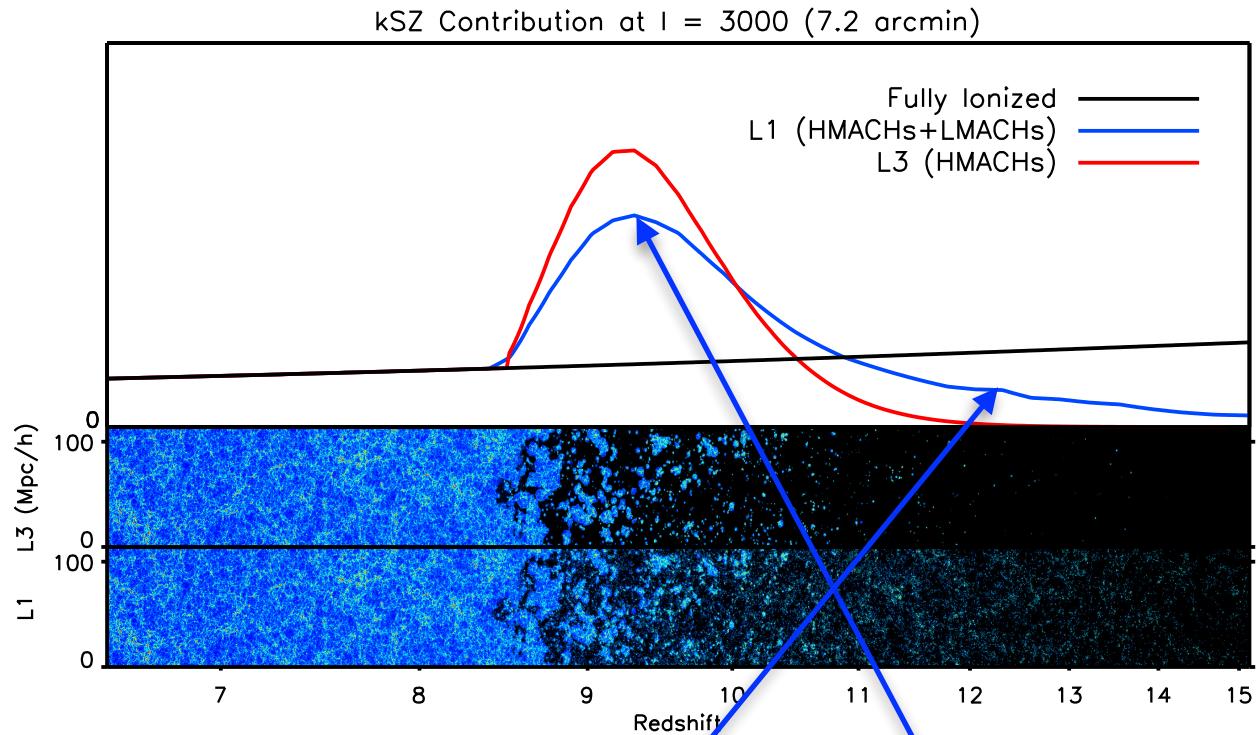


Low-mass sources added



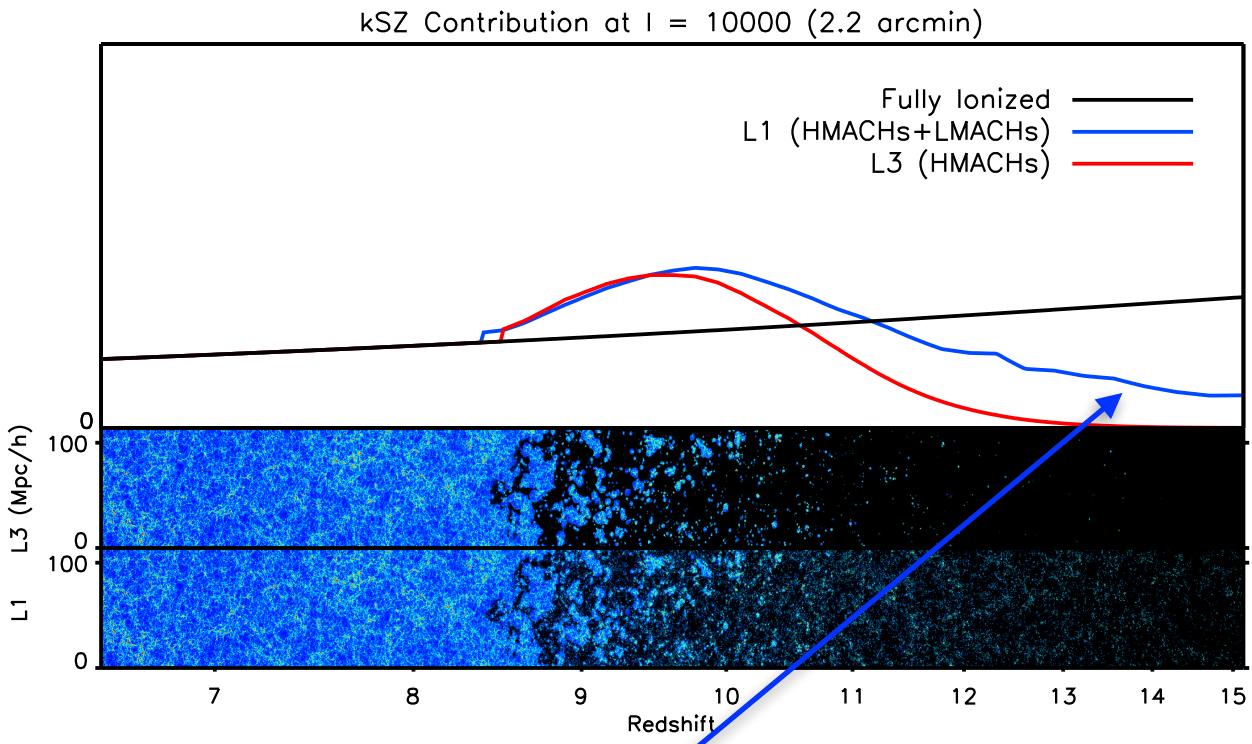
(Park et al. 2013)

# 1-1) Models with and without Low-mass Galaxies ( $l = 3000$ )



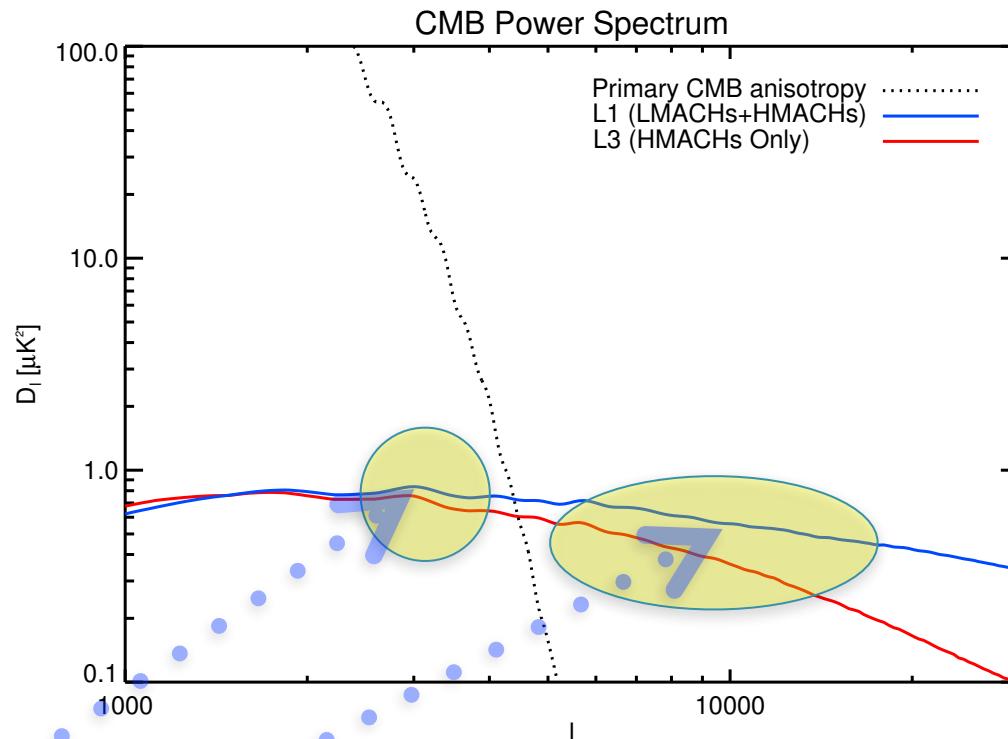
Including low-mass galaxies give a lower signal at late time and higher signal at early time. Both effects nearly cancel each other out to give a similar total signal,  $0.9 \mu\text{K}^2$ .

# 1-1) Models with and without Low-mass Galaxies ( $I = 10000$ )



**The low-mass galaxies have a larger impact in smaller angular scales.**

# Result 1-1) Impact of low-mass galaxies on the kSZ Power Spectrum



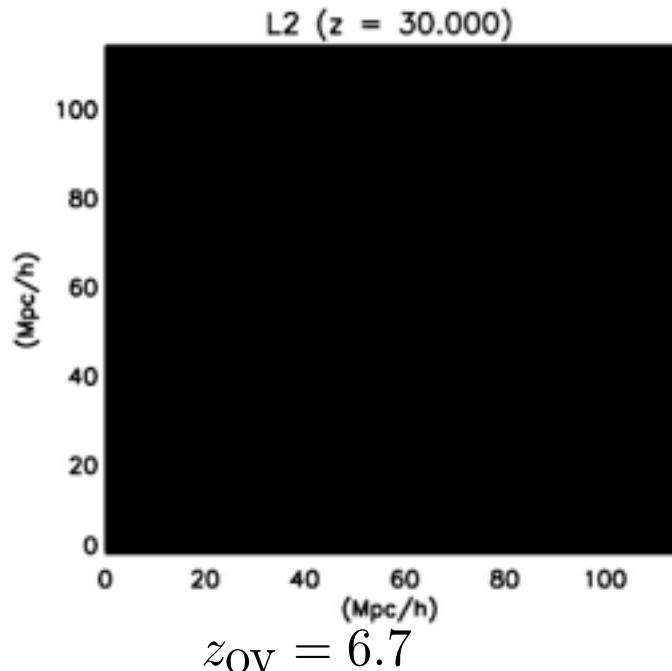
- At  $l = 3000$  ( $\sim 3$  arcmin), adding low-mass galaxies make kSZ signal only 5% larger.
- At larger  $l$ 's, the difference is larger.

# 1-2) Models with and without Mihihalo Galaxies

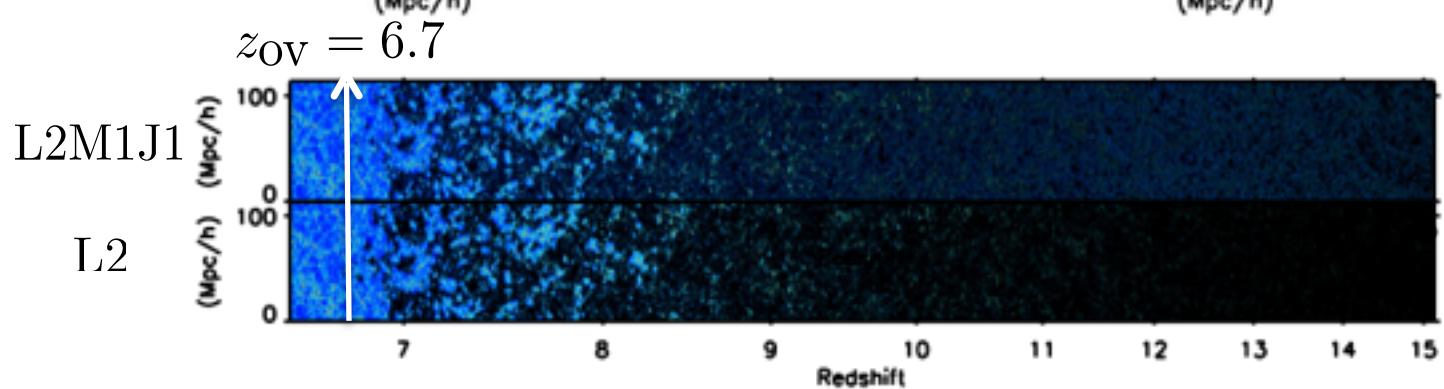
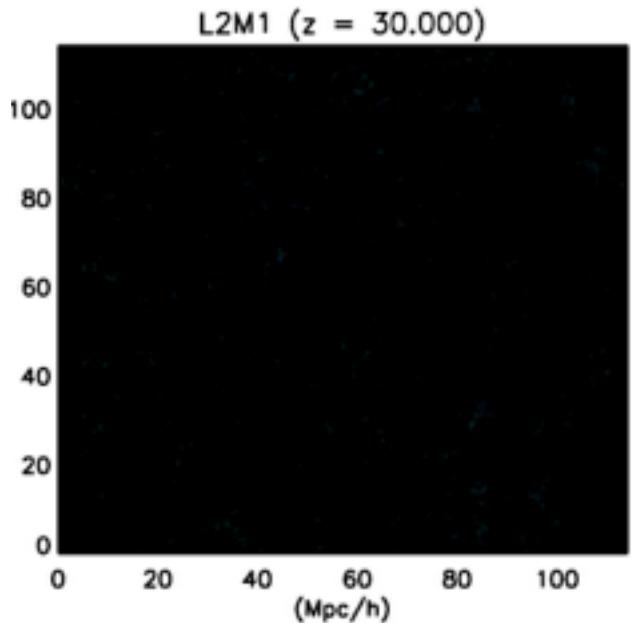
- **High-mass Atomic Cooling Halos:  $M > 10^9 M_\odot$**   
**Never suppressed.**
- **Low-mass Atomic Cooling Halos:  $10^8 M_\odot < M < 10^9 M_\odot$**   
**Suppressed when surrounding region is ionized.**
- **Minihalos:  $M < 10^8 M_\odot$**   
**Suppressed when surrounding region is ionized or exposed to H<sub>2</sub> dissociating radiation.**

# 1-2) Models with and without Mihihalo Galaxies

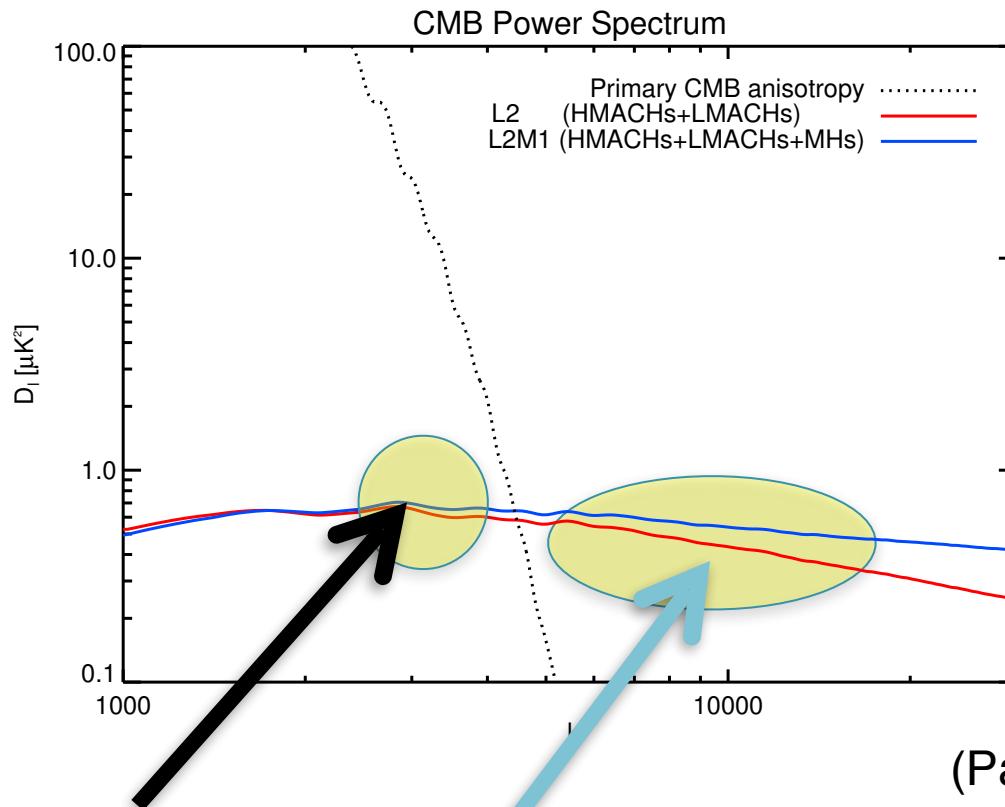
HMACs and LMACHs



HMACs, LMACHs and MHs

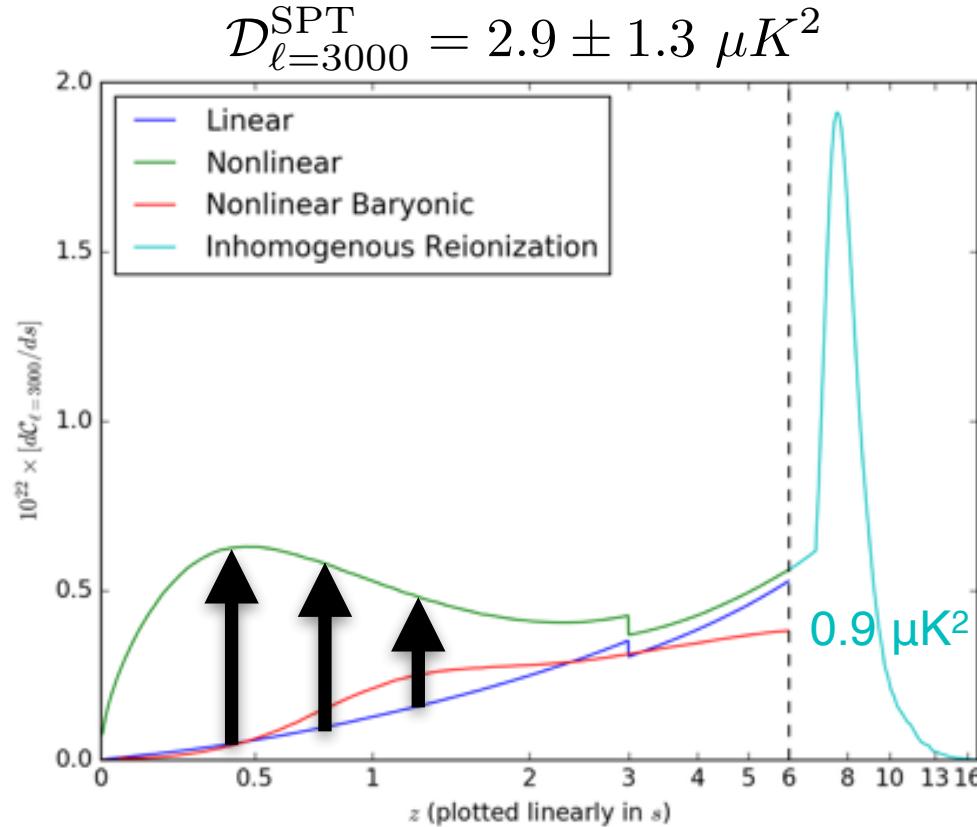


# Result 1-2) Impact of Minihalo galaxies on kSZ Power Spectrum



- At  $l = 3000$ , adding MHs make kSZ signal only 3% larger.
- At larger l's, the difference is larger.

## Part 2) Impact of Nonlinear Growth of Structure on the Post-reionization Signal



$$\mathbf{q} = X(1 + \delta)\mathbf{v} = (1 + \delta_{nl})\mathbf{v}_{nl}$$

**Ionized fraction (X) is one everywhere. But, density ( $\delta$ ) and velocity ( $v$ ) are in the nonlinear regime.**

# Expressing $P_{q,\text{perp}}$ in terms of density( $\delta$ ) and velocity( $\mathbf{v}$ )

In the post-reionization epoch, IGM is fully ionized to a good approximation. Therefore, we can set  $X = 1$ .

$$\mathbf{q} = \mathbf{v}(1 + \delta) = \mathbf{v} + \mathbf{v}\delta$$

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$$\tilde{\mathbf{q}}(\mathbf{k}) = \tilde{\mathbf{v}}(\mathbf{k}) + \int \frac{d^3 k'}{(2\pi)^3} \tilde{\mathbf{v}}(\mathbf{k}') \tilde{\delta}(\mathbf{k} - \mathbf{k}')$$

---

$$\tilde{\mathbf{q}}_{\perp}(\mathbf{k}) = \tilde{\mathbf{v}}_{\perp}(\mathbf{k}) + \left( \int \frac{d^3 k'}{(2\pi)^3} \tilde{\mathbf{v}}(\mathbf{k}') \tilde{\delta}(\mathbf{k} - \mathbf{k}') \right)_{\perp}$$

Since the transverse mode is negligibly small for velocity, the transverse momentum field has only a 2<sup>nd</sup> order term.

# $P_{q,\text{perp}}$ in terms of $\delta$ and $\mathbf{v}$

$$(2\pi)^3 P_{q\perp}(\mathbf{k}_1) \delta_D(\mathbf{k}_1 + \mathbf{k}_2) = \langle \tilde{\mathbf{q}}_\perp(\mathbf{k}_1) \tilde{\mathbf{q}}_\perp(\mathbf{k}_2) \rangle$$

$$= \left\langle \int \frac{d^3 k'}{(2\pi)^3} \tilde{\mathbf{v}}_\perp(\mathbf{k}') \delta(\mathbf{k}_1 - \mathbf{k}') \cdot \int \frac{d^3 k'}{(2\pi)^3} \tilde{\mathbf{v}}_\perp(\mathbf{k}') \delta(\mathbf{k}_2 - \mathbf{k}') \right\rangle$$

---

$$P_{q\perp} = \sum_{i=1}^3 \left[ 1 - (\hat{k}^i)^2 \right] P_q^{ii}$$

where

$$(2\pi)^3 P_q^{ij}(k) \delta_D(\mathbf{k}_1 + \mathbf{k}_2)$$
$$= \int \frac{d^3 k'}{(2\pi)^3} \int \frac{d^3 k''}{(2\pi)^3} \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^i(\mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle$$

# $P_{q,\text{perp}}$ in terms of $\delta$ and $v$

Using Wick's theorem

( $\langle ABCD \rangle = \langle AB \rangle \langle CD \rangle + \langle AC \rangle \langle BD \rangle + \langle AD \rangle \langle BC \rangle$ ),  
the 4th order bracket can be decomposed into pair brackets.

$$\begin{aligned} & \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &= \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \right\rangle \left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &+ \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \right\rangle \left\langle \tilde{v}^i(\mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &+ \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle \left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^i(\mathbf{k}') \right\rangle \end{aligned}$$

# $P_{q,\text{perp}}$ in terms of $\delta$ and $v$

$$\begin{aligned} & \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &= \cancel{\left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \right\rangle} \cancel{\left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle} \end{aligned}$$

$$\begin{aligned} & + \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \right\rangle \langle \tilde{v}^i(\mathbf{k}') \tilde{v}^j(\mathbf{k}'') \rangle \\ & + \cancel{\left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle} \cancel{\left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^i(\mathbf{k}') \right\rangle} \end{aligned}$$

$$P_{q_\perp} = \int \frac{d^3 k'}{(2\pi)^3} \left( 1 - \mu'^2 \right)$$

$$\left[ P_{\delta\delta}(|\mathbf{k} - \mathbf{k}'|) P_{vv}(\mathbf{k}') - \frac{k'}{|\mathbf{k} - \mathbf{k}'|} P_{\delta v}(|\mathbf{k} - \mathbf{k}'|) P_{\delta v}(\mathbf{k}') \right]$$

$$(\mu' \equiv \hat{k} \cdot \hat{k}')$$

(Ma & Fry 2002)

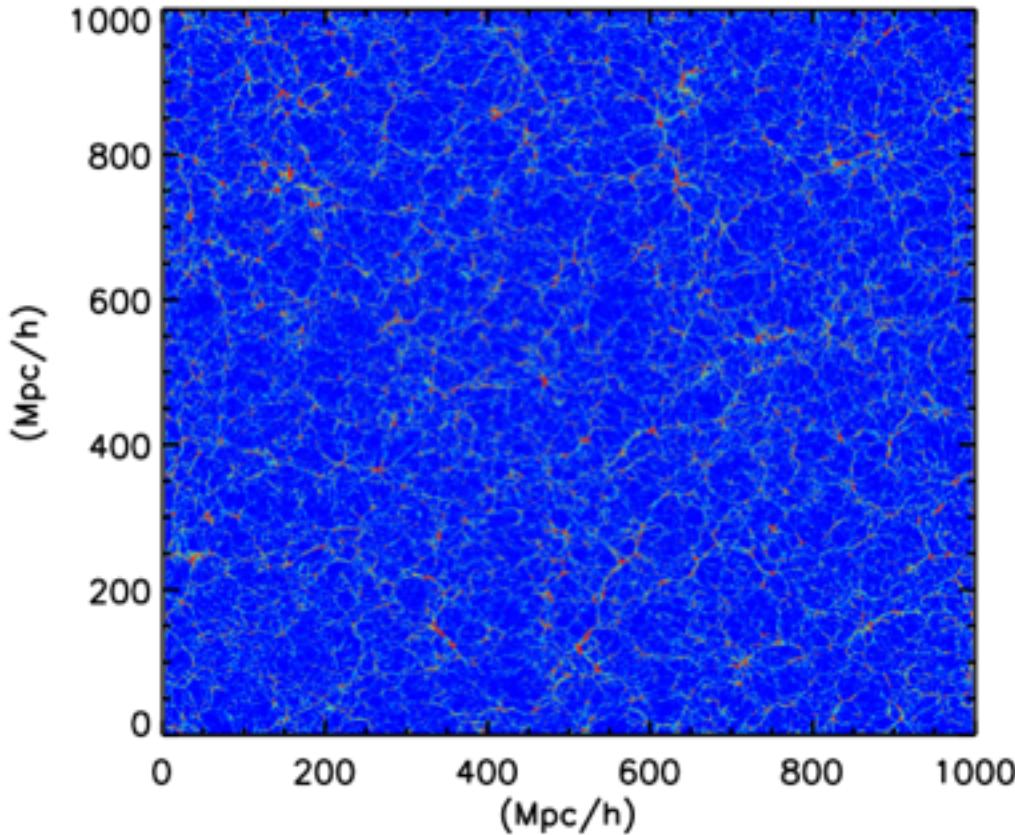
# $P_{q,\text{perp}}$ in terms of $\delta$ and $v$

$$P_{q\perp} = \int \frac{d^3 k'}{(2\pi)^3} \left( 1 - \mu'^2 \right) \left[ P_{\delta\delta}(|\mathbf{k} - \mathbf{k}'|) P_{vv}(\mathbf{k}') - \frac{k'}{|\mathbf{k} - \mathbf{k}'|} P_{\delta v}(|\mathbf{k} - \mathbf{k}'|) P_{\delta v}(\mathbf{k}') \right]$$
$$(\mu' \equiv \hat{\mathbf{k}} \cdot \hat{\mathbf{k}}')$$

(Ma & Fry 2002)

**Does this expression really hold in the nonlinear regime?**

# Testing the analytic expression of $P_{q,\text{perp}}$



**Code : CubeP3M**

**Size : 1 Gpc/h**

**# of ptls :  $3456^3$**

We use *N*-body simulation data to test the non-linear expression

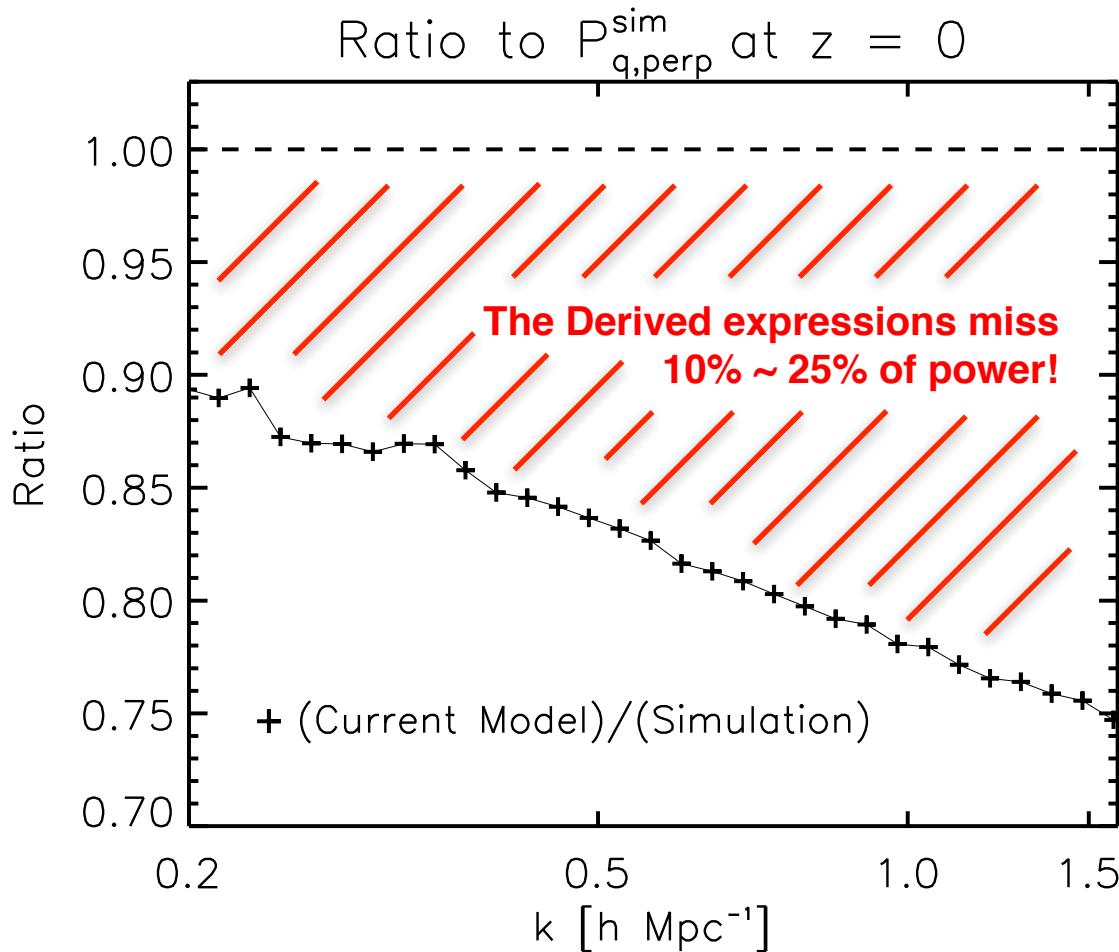
# $P_{q,\text{perp}}$ in terms of $\delta$ and $v$

$$P_{q\perp} = \int \frac{d^3 k'}{(2\pi)^3} \left( 1 - \mu'^2 \right) \left[ P_{\delta\delta}(|\mathbf{k} - \mathbf{k}'|) P_{vv}(\mathbf{k}') - \frac{k'}{|\mathbf{k} - \mathbf{k}'|} P_{\delta v}(|\mathbf{k} - \mathbf{k}'|) P_{\delta v}(\mathbf{k}') \right]$$
$$(\mu' \equiv \hat{\mathbf{k}} \cdot \hat{\mathbf{k}}')$$

(Ma & Fry 2002)

**Does this expression really hold in the nonlinear regime?**

# Testing the analytic expression of $P_{q,\text{perp}}$



# What is wrong with the analytic expression?

Wick's theorem works only for gaussian fields. In the nonlinear regime where density and velocity becomes non-gaussian, we need a correction term called “Connected Moment”.

$$\begin{aligned} & \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &= \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^i(\mathbf{k}') \right\rangle \left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &+ \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \right\rangle \left\langle \tilde{v}^i(\mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle \\ &+ \left\langle \tilde{\delta}(\mathbf{k}_1 - \mathbf{k}') \tilde{v}^j(\mathbf{k}'') \right\rangle \left\langle \tilde{\delta}(\mathbf{k}_2 - \mathbf{k}'') \tilde{v}^i(\mathbf{k}') \right\rangle \\ &+ (\text{Connected Moment}) \end{aligned}$$

# Perturbation Theory Expansion

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$$\delta = \delta^{(1)} + \delta^{(2)} + \dots$$

$$\delta^{(2)}(\mathbf{k}) = \int \frac{d^3 k_a}{(2\pi)^3} \int d^3 k_b \ \delta^{(1)}(\mathbf{k}_a) \delta^{(1)}(\mathbf{k}_b) F_2(\mathbf{k}_a, \mathbf{k}_b) \delta_D(\mathbf{k} - \mathbf{k}_a - \mathbf{k}_b)$$

$$F_2(\mathbf{q}_1, \mathbf{q}_2) = \frac{5}{7} + \frac{1}{2} \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1 q_2} + \frac{2}{7} \left( \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1 q_2} \right)^2$$

---

$$\theta \equiv \nabla \cdot \mathbf{v} = \theta^{(1)} + \theta^{(2)} + \dots$$

$$\theta^{(2)}(\mathbf{k}) = - \int \frac{d^3 k_a}{(2\pi)^3} \int d^3 k_b \ \delta^{(1)}(\mathbf{k}_a) \delta^{(1)}(\mathbf{k}_b) G_2(\mathbf{k}_a, \mathbf{k}_b) \delta_D(\mathbf{k} - \mathbf{k}_a - \mathbf{k}_b)$$

$$G_2(\mathbf{q}_1, \mathbf{q}_2) = \frac{3}{7} + \frac{1}{2} \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1 q_2} + \frac{4}{7} \left( \frac{\mathbf{q}_1 \cdot \mathbf{q}_2}{q_1 q_2} \right)^2$$

---

**Nonlinear terms can be expressed in terms of linear terms using perturbation theory**

# P<sub>q</sub>: Next-to-leading order

$$\left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\mathbf{v}}(\mathbf{k}_3)\tilde{\mathbf{v}}(\mathbf{k}_4) \right\rangle = -\dot{a}^2 \left( \frac{\hat{\mathbf{k}}_3}{k_3} \right) \left( \frac{\hat{\mathbf{k}}_4}{k_4} \right) \left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\theta}(\mathbf{k}_3)\tilde{\theta}(\mathbf{k}_4) \right\rangle$$

$$\delta = \delta^{(1)} + \delta^{(2)} + \dots \quad \theta \equiv i\nabla \cdot \mathbf{v} = \theta^{(1)} + \theta^{(2)} + \dots$$

Next-to-leading order of  $\left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\theta}(\mathbf{k}_3)\tilde{\theta}(\mathbf{k}_4) \right\rangle$  is

---

$$\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle$$

+

$$\left\langle \tilde{\delta}^{(3)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(3)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(3)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(3)}(\mathbf{k}_4) \right\rangle$$

# Connected Moment of $P_{q,\text{perp}}$

## Next-to-leading order

$$\begin{aligned}
 & \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(2)}(\mathbf{k}_2) \tilde{\theta}^{(2)}(\mathbf{k}_3) \tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle \\
 = & \int \frac{d^3 \mathbf{k}_{2,a}}{(2\pi)^3} \int d^3 \mathbf{k}_{2,b} \int \frac{d^3 \mathbf{k}_{3,a}}{(2\pi)^3} \int d^3 \mathbf{k}_{3,b} \\
 & F_2^{(s)}(\mathbf{k}_{2,a}, \mathbf{k}_{2,b}) G_2^{(s)}(\mathbf{k}_{3,a}, \mathbf{k}_{3,b}) \delta_D(\mathbf{k}_2 - \mathbf{k}_{2,a} - \mathbf{k}_{2,b}) \delta_D(\mathbf{k}_3 - \mathbf{k}_{3,a} - \mathbf{k}_{3,b}) \\
 & \overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(1)}(\mathbf{k}_{2,a}) \tilde{\delta}^{(1)}(\mathbf{k}_{2,b}) \tilde{\delta}^{(1)}(\mathbf{k}_{3,a}) \tilde{\delta}^{(1)}(\mathbf{k}_{3,b}) \tilde{\delta}^{(1)}(\mathbf{k}_4) \right\rangle} \\
 = & \int \frac{d^3 \mathbf{k}_{2,a}}{(2\pi)^3} \int d^3 \mathbf{k}_{2,b} \int \frac{d^3 \mathbf{k}_{3,a}}{(2\pi)^3} \int d^3 \mathbf{k}_{3,b} \\
 & F_2^{(s)}(\mathbf{k}_{2,a}, \mathbf{k}_{2,b}) G_2^{(s)}(\mathbf{k}_{3,a}, \mathbf{k}_{3,b}) \delta_D(\mathbf{k}_2 - \mathbf{k}_{2,a} - \mathbf{k}_{2,b}) \delta_D(\mathbf{k}_3 - \mathbf{k}_{3,a} - \mathbf{k}_{3,b}) \\
 & \overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(1)}(\mathbf{k}_{2,a}) \right\rangle \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_{2,b}) \tilde{\delta}^{(1)}(\mathbf{k}_{3,a}) \right\rangle \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_{3,b}) \tilde{\delta}^{(1)}(\mathbf{k}_4) \right\rangle}
 \end{aligned}$$

15 terms in total

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$= \int \frac{d^3 k_{2,a}}{(2\pi)^3} \int d^3 k_{2,b} \int \frac{d^3 k_{3,a}}{(2\pi)^3} \int d^3 k_{3,b} \\ F_2^{(s)}(k_{2,a}, k_{2,b}) G_2^{(s)}(k_{3,a}, k_{3,b}) \delta_D(k_2 - k_{2,a} - k_{2,b}) \delta_D(k_3 - k_{3,a} - k_{3,b})$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$\begin{aligned}
& \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(2)}(\mathbf{k}_2) \tilde{\theta}^{(2)}(\mathbf{k}_3) \tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle_{uc} \\
&= \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(2)}(\mathbf{k}_2) \right\rangle \left\langle \tilde{\theta}^{(2)}(\mathbf{k}_3) \tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle \\
&+ \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(2)}(\mathbf{k}_3) \right\rangle \left\langle \tilde{\theta}^{(2)}(\mathbf{k}_2) \tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle \\
&+ \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(1)}(\mathbf{k}_4) \right\rangle \left\langle \tilde{\theta}^{(2)}(\mathbf{k}_2) \tilde{\theta}^{(2)}(\mathbf{k}_3) \right\rangle
\end{aligned}$$

$$= \int \frac{d^3 k_{2,a}}{(2\pi)^3} \int d^3 k_{2,b} \int \frac{d^3 k_{3,a}}{(2\pi)^3} \int d^3 k_{3,b} \\ F_2^{(s)}(k_{2,a}, k_{2,b}) G_2^{(s)}(k_{3,a}, k_{3,b}) \delta_D(k_2 - k_{2,a} - k_{2,b}) \delta_D(k_3 - k_{3,a} - k_{3,b})$$

$$\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle$$

$$= \left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1) \tilde{\delta}^{(1)}(\mathbf{k}_4) \right\rangle \left\langle \tilde{\theta}^{(2)}(\mathbf{k}_2) \tilde{\theta}^{(2)}(\mathbf{k}_3) \right\rangle$$

$$= \int \frac{d^3 k_{2,a}}{(2\pi)^3} \int d^3 k_{2,b} \int \frac{d^3 k_{3,a}}{(2\pi)^3} \int d^3 k_{3,b} \\ F_2^{(s)}(k_{2,a}, k_{2,b}) G_2^{(s)}(k_{3,a}, k_{3,b}) \delta_D(k_2 - k_{2,a} - k_{2,b}) \delta_D(k_3 - k_{3,a} - k_{3,b})$$

# → Connected Moment

# Unconnected Moment

# P<sub>q</sub>: Next-to-leading order

$$\left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\mathbf{v}}(\mathbf{k}_3)\tilde{\mathbf{v}}(\mathbf{k}_4) \right\rangle = -\dot{a}^2 \left( \frac{\hat{\mathbf{k}}_3}{k_3} \right) \left( \frac{\hat{\mathbf{k}}_4}{k_4} \right) \left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\theta}(\mathbf{k}_3)\tilde{\theta}(\mathbf{k}_4) \right\rangle$$

$$\delta = \delta^{(1)} + \delta^{(2)} + \dots \quad \theta \equiv i\nabla \cdot \mathbf{v} = \theta^{(1)} + \theta^{(2)} + \dots$$

Next-to-leading order of  $\left\langle \tilde{\delta}(\mathbf{k}_1)\tilde{\delta}(\mathbf{k}_2)\tilde{\theta}(\mathbf{k}_3)\tilde{\theta}(\mathbf{k}_4) \right\rangle$  is

---

$$\overline{\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(2)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(2)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(2)}(\mathbf{k}_3)\tilde{\theta}^{(2)}(\mathbf{k}_4) \right\rangle}$$

+

$$\overline{\left\langle \tilde{\delta}^{(3)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(3)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(3)}(\mathbf{k}_3)\tilde{\theta}^{(1)}(\mathbf{k}_4) \right\rangle}$$

$$\overline{\left\langle \tilde{\delta}^{(1)}(\mathbf{k}_1)\tilde{\delta}^{(1)}(\mathbf{k}_2)\tilde{\theta}^{(1)}(\mathbf{k}_3)\tilde{\theta}^{(3)}(\mathbf{k}_4) \right\rangle}$$

# Connected Moment of $P_{q,\text{perp}}$

## Next-to-leading order

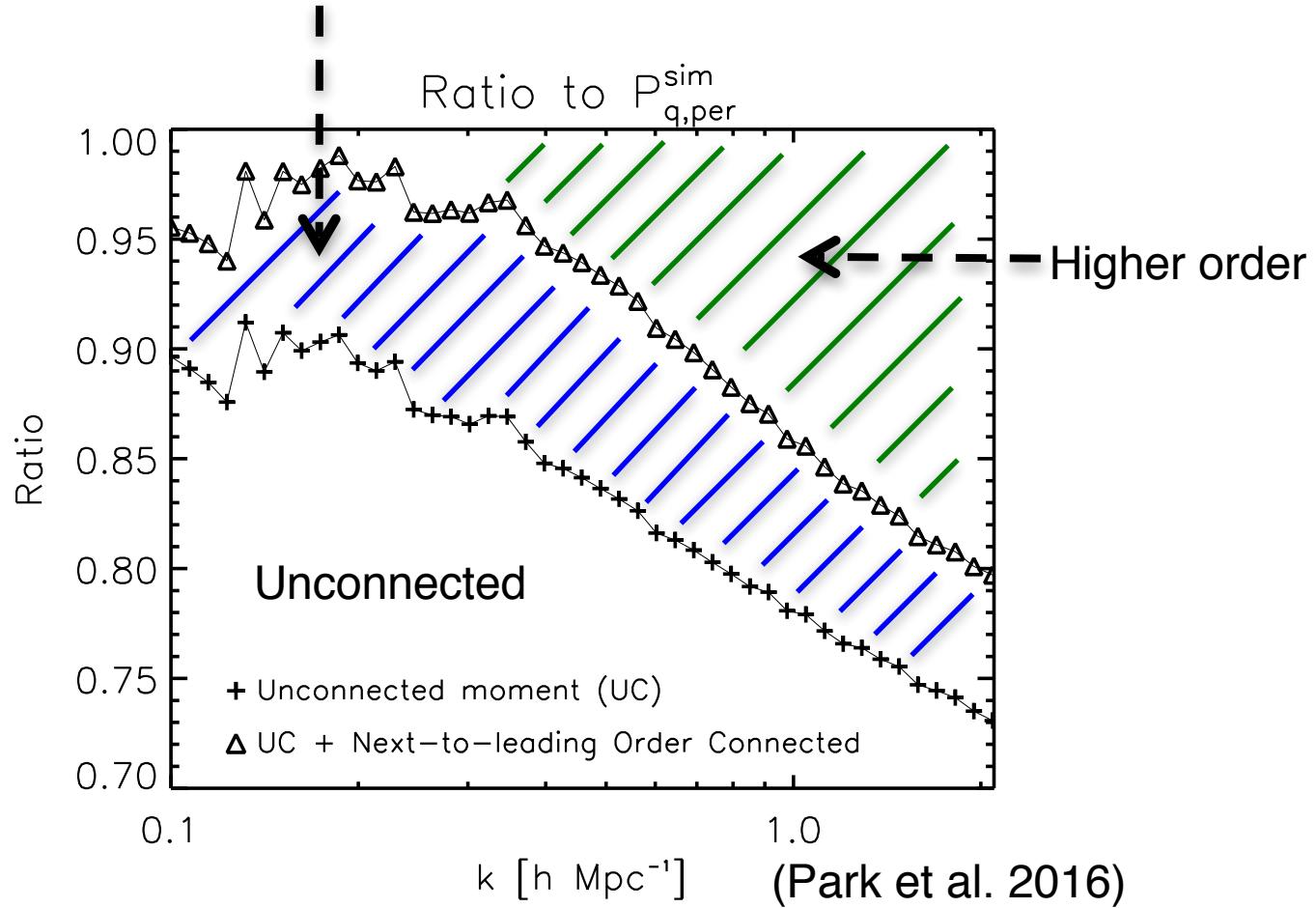
$$\begin{aligned}
 P_{q\perp,c} = & - \int \frac{d^3 k'}{(2\pi)^3} \int \frac{d^3 k''}{(2\pi)^3} \left[ \frac{1 - (\hat{k} \cdot \hat{k}')(\hat{k} \cdot \hat{k}'')}{k' k''} \right] \\
 & [6[F_3^{(s)}(k + k'', -k', -k'')P(k + k'')P(k')P(k'') \\
 & + F_3^{(s)}(k - k', k', k'')P(k - k')P(k')P(k'') \\
 & + G_3^{(s)}(k - k', -k - k'', k'')P(k - k')P(k + k'')P(k'') \\
 & + G_3^{(s)}(k - k', -k - k'', k')P(k - k')P(k + k'')P(k')] \\
 & + 4[F_2^{(s)}(k - k', k' + k'')G_2^{(s)}(k' + k'', -k'')P(k - k')P(k' + k'')P(k'') \\
 & + F_2^{(s)}(k - k', k' + k'')G_2^{(s)}(-k' - k'', k')P(k - k')P(k' + k'')P(k') \\
 & + G_2^{(s)}(k, -k + k')F_2^{(s)}(k, k'')P(k - k')P(k)P(k'') \\
 & + G_2^{(s)}(k, -k + k')G_2^{(s)}(k, -k - k'')P(k - k')P(k)P(k + k'') \\
 & + G_2^{(s)}(k - k' + k'', -k + k')F_2^{(s)}(k - k' + k'', k')P(k - k')P(k - k' + k'')P(k') \\
 & + G_2^{(s)}(k - k' + k'', -k + k')G_2^{(s)}(k - k' + k'', -k - k'')P(k - k')P(k - k' + k'')P(k + k'') \\
 & + F_2^{(s)}(-k' - k'', k + k'')G_2^{(s)}(-k' - k'', k'')P(k + k'')P(k' + k'')P(k'') \\
 & + F_2^{(s)}(-k' - k'', k + k'')G_2^{(s)}(-k' - k'', k')P(k + k'')P(k' + k'')P(k') \\
 & + G_2^{(s)}(-k - k'' + k', k + k'')F_2^{(s)}(-k - k'' + k', k'')P(k + k'')P(k - k' + k'')P(k'') \\
 & + G_2^{(s)}(-k, k + k'')F_2^{(s)}(-k, k')P(k + k'')P(k)P(k') \\
 & + F_2^{(s)}(k, -k')F_2^{(s)}(k, k'')P(k)P(k')P(k'') \\
 & + F_2^{(s)}(k - k' + k'', k')F_2^{(s)}(k - k' + k'', -k'')P(k')P(k - k' + k'')P(k'')]]
 \end{aligned}$$

We have derived the connected term in the next-to-leading order.

(Park et al. 2016)

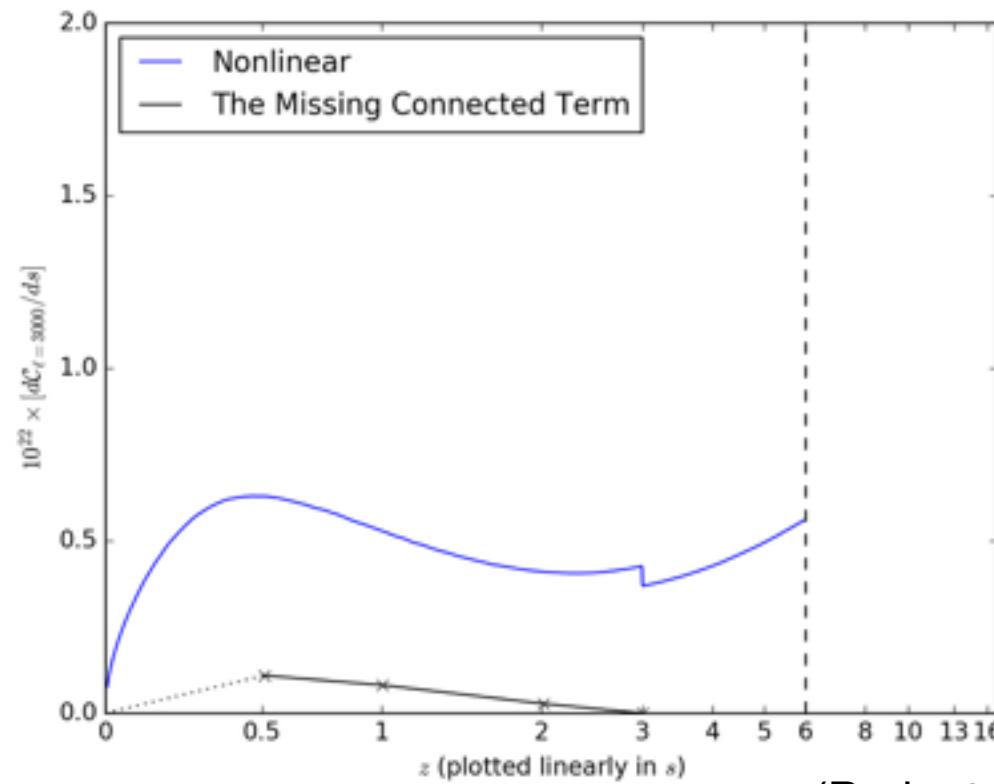
# Connected Moment of $P_{q,\text{perp}}$

## Next-to-leading order



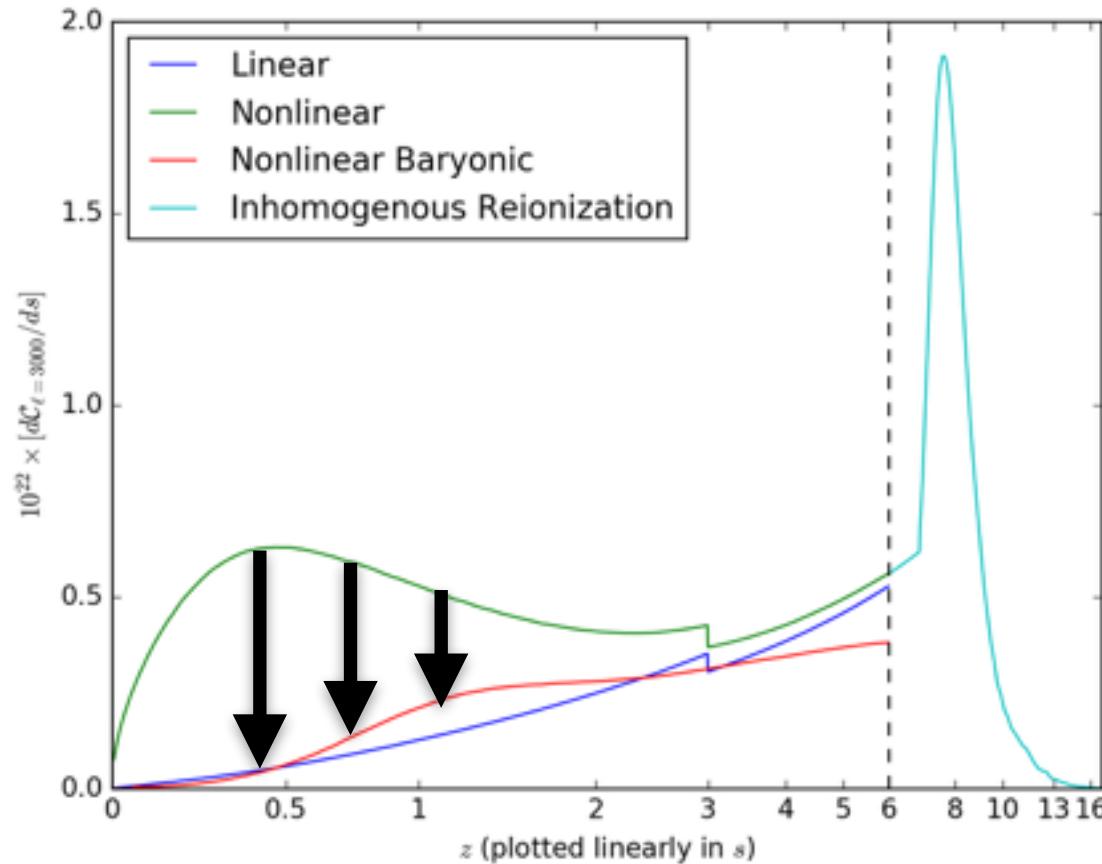
Next-to-leading order explains the deficit up to  $k \sim 0.3 \text{ h/Mpc}$ !  
At  $k > 0.3 \text{ h/Mpc}$ , perhaps higher-order terms dominates.

# Final Result of Part 2



The connected term add ~10% extra to the post-reionization kSZ signal compared to the previous expression.

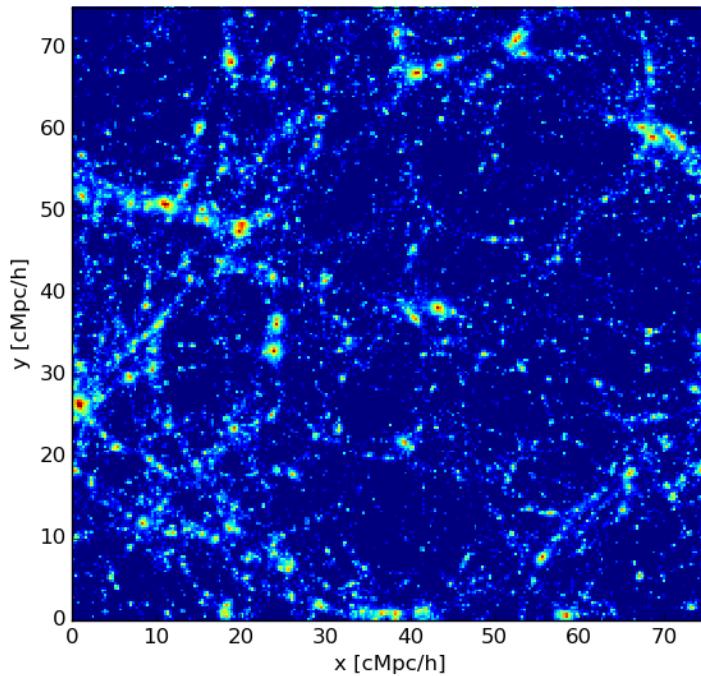
# Part 3) Effect of Baryonic Physics in the kSZ signal



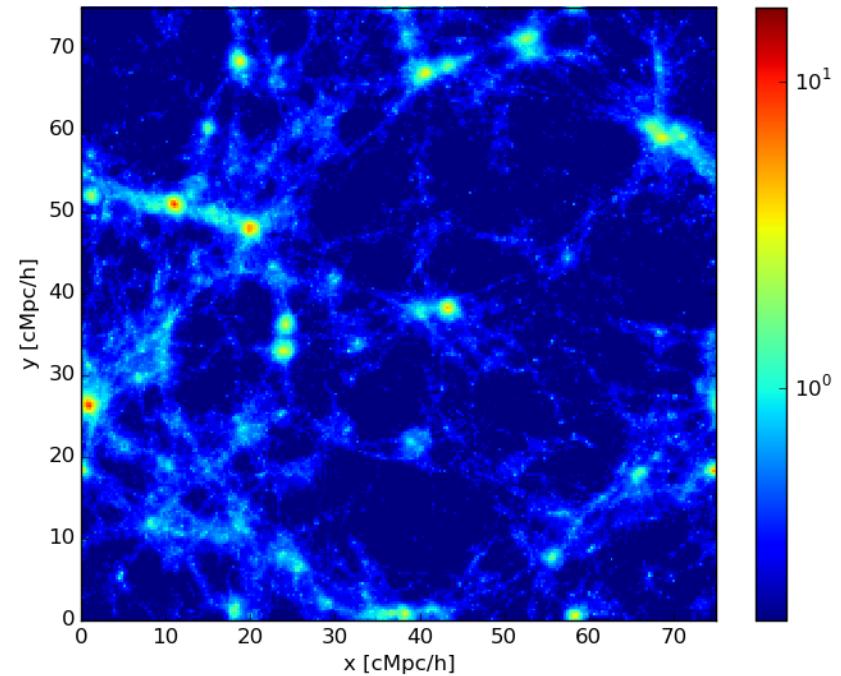
At low redshift, offset between dark matter and free electron densities due to baryonic physics becomes an important factor.

# Part 3) Impact of Baryonic Physics on the kSZ signal

Dark matter



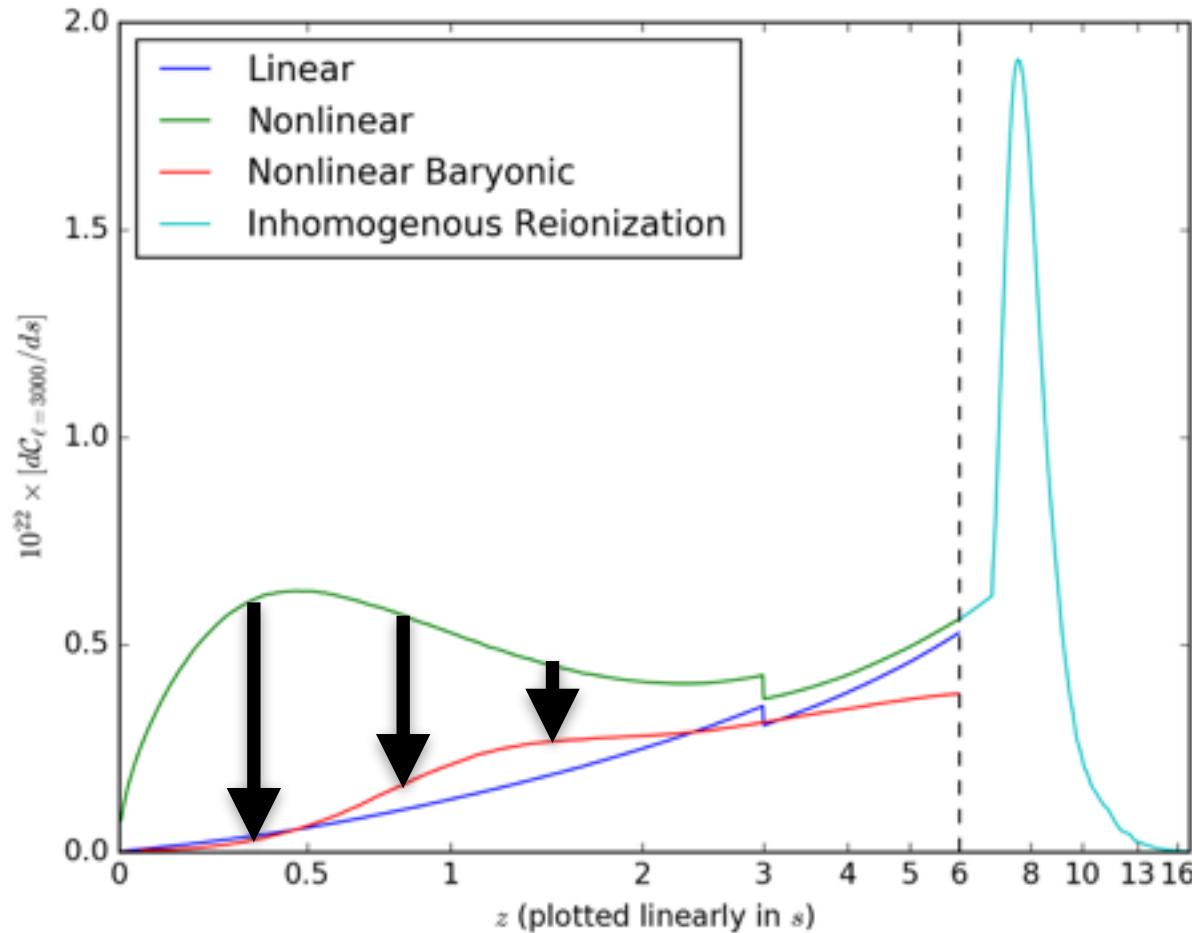
Gas



**Two things happen to the baryonic density field at low  $z$ :**

- 1) Star formation locks some gas into stars.
- 2) Star formation feedback wipes out density structures in gas.

# Part 3) Effect of Baryonic Physics in the kSZ signal



$$\mathbf{q} = X(1 + \delta)\mathbf{v} = (1 + \delta_{nl,e})\mathbf{v}_{nl,e}$$

**Both effects act toward suppressing the signal.**

# How to model the baryonic effect in the kSZ signal

The first step is to use the electron density power spectrum in the equation below.

$$P_{q\perp} = \int \frac{d^3 k'}{(2\pi)^3} \left( 1 - \mu'^2 \right) \\ \left[ P_{\delta\delta}(|\mathbf{k} - \mathbf{k}'|) P_{vv}(\mathbf{k}') - \frac{k'}{|\mathbf{k} - \mathbf{k}'|} P_{\delta v}(|\mathbf{k} - \mathbf{k}'|) P_{\delta v}(\mathbf{k}') \right]$$

---

Assume the baryonic density field is related to the dark matter density field by a window function,  $W$ .

$$\tilde{\delta}_b(k) = W(k) \tilde{\delta}_{dm}(k) \quad \longrightarrow \quad P_{\delta\delta}^b(k) = W^2(k) P_{\delta\delta}^{dm}(k)$$
$$P_{\delta v}^b(k) = W(k) P_{\delta v}^{dm}(k)$$

# Illustris Simulation

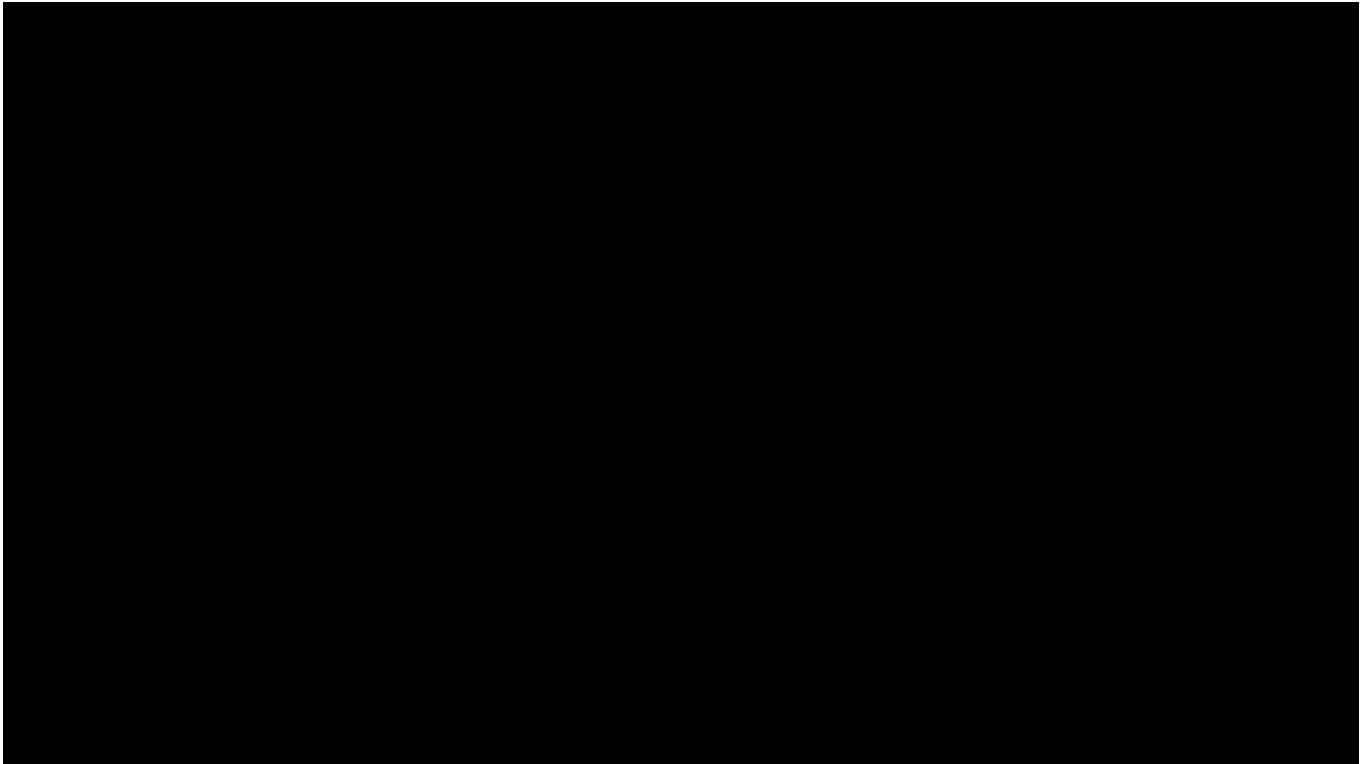
## Physics included (in addition to gravity)

1. Gas Cooling
2. Star Formation
3. Stellar Evolution
4. Stellar Radiation Feedback
5. Supernova Feedback
6. Blackhole Feedback

| Simulation Name  | $L_{\text{box}} [\text{Mpc}]$ | $N_{\text{DM}}$ | $m_{\text{DM}} [M_{\odot}]$ | $m_{\text{gas}} [M_{\odot}]$ |
|------------------|-------------------------------|-----------------|-----------------------------|------------------------------|
| Illustris-1      | 106.5                         | $1820^3$        | $6.3 \times 10^6$           | $1.3 \times 10^6$            |
| Illustris-1-Dark | 106.5                         | $1820^3$        | $7.5 \times 10^6$           | 0                            |
| Illustris-2      | 106.5                         | $910^3$         | $5.0 \times 10^7$           | $1.0 \times 10^7$            |
| Illustris-2-Dark | 106.5                         | $910^3$         | $6.0 \times 10^7$           | 0                            |
| Illustris-3      | 106.5                         | $455^3$         | $4.0 \times 10^8$           | $8.1 \times 10^7$            |
| Illustris-3-Dark | 106.5                         | $455^3$         | $4.8 \times 10^8$           | 0                            |

Illustris is a high resolution simulation that includes a list of crucial baryonic physics that allows to create present day galaxies.

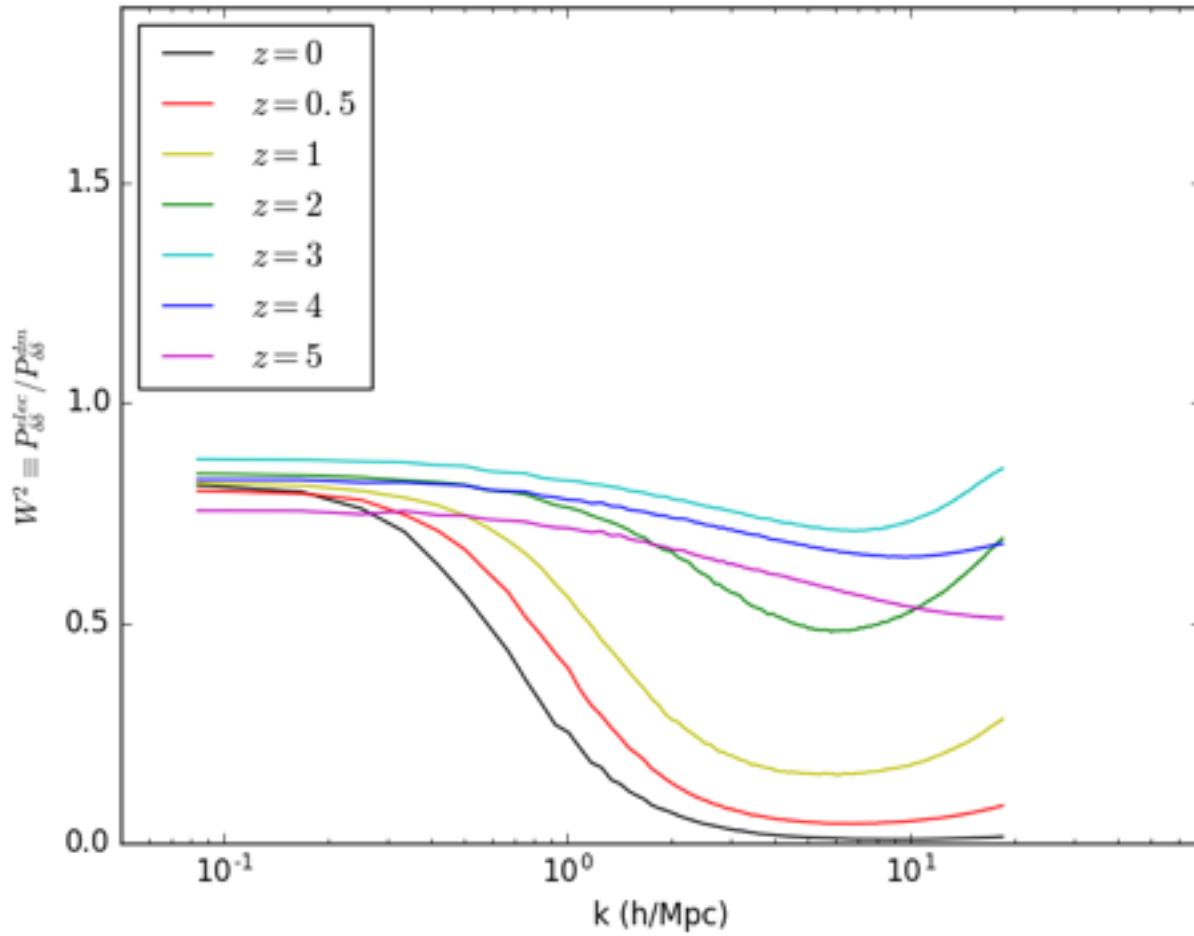
# Illustris Simulation



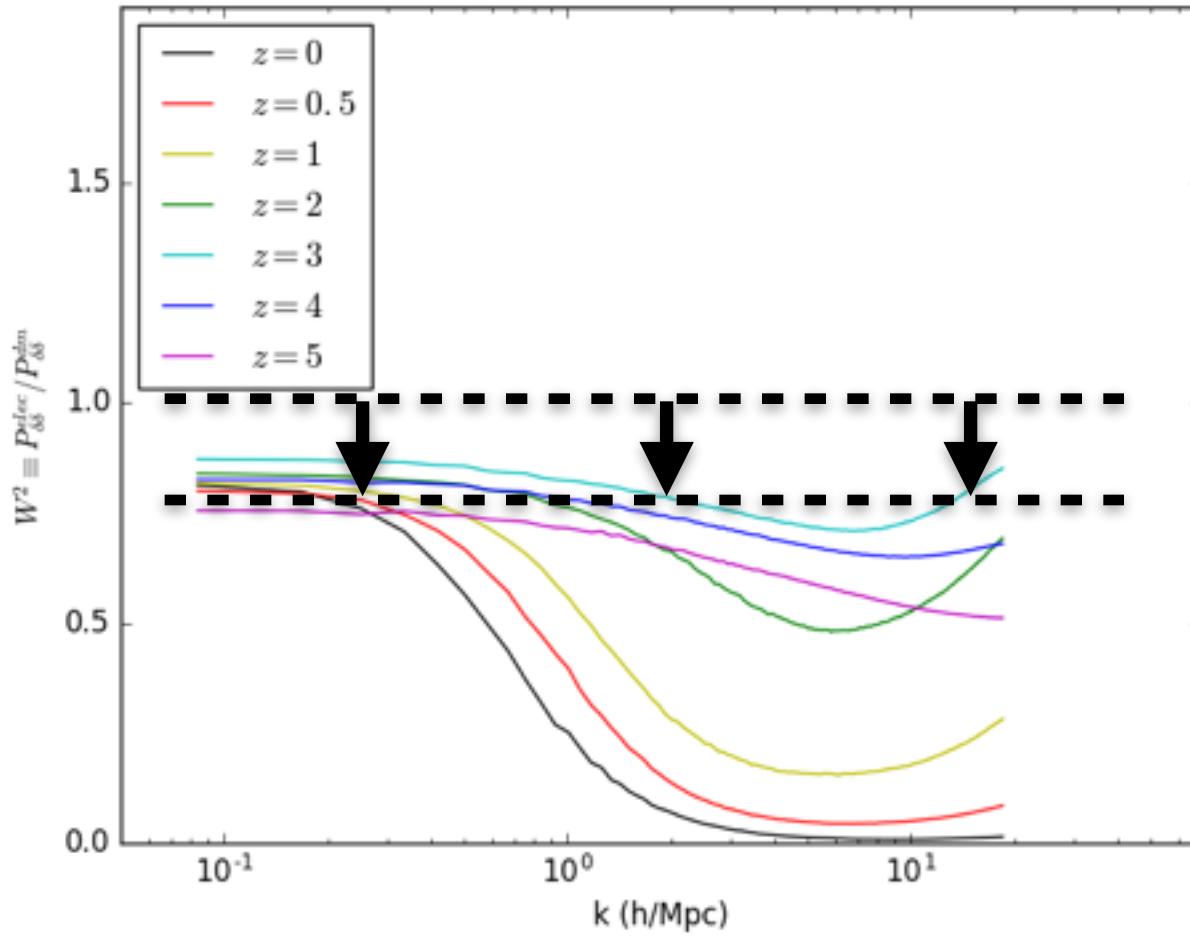
Video from <http://www.illustris-project.org/media/>

**Illustris has a fairly large size (75 Mpc/h) with baryonic physics (star-formation & feedback) needed for kSZ calculation.**

# Window function from Illustris



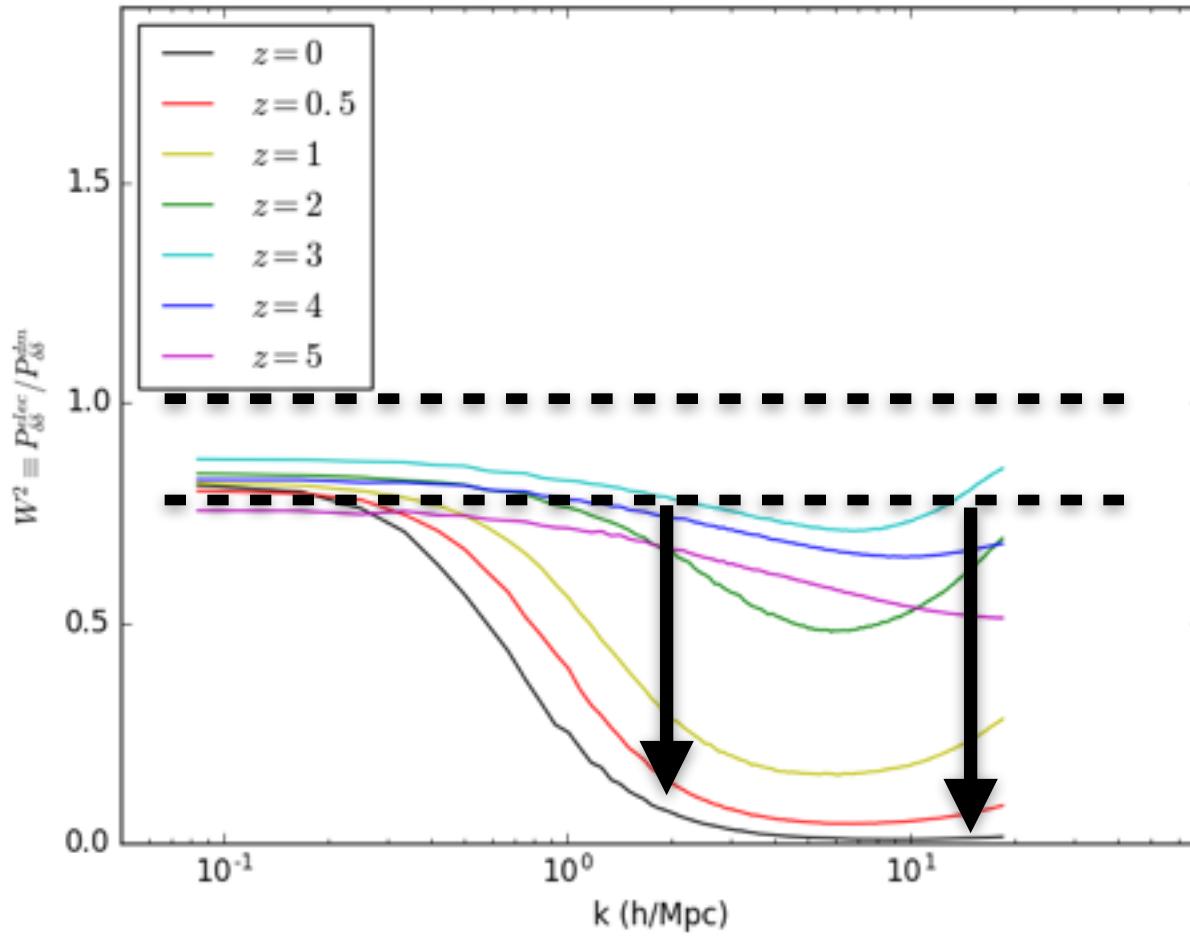
# Window function from Illustris



Suppression at all scales due to gas locked into stellar products.

(Park et al. in prep.)

# Window function from Illustris

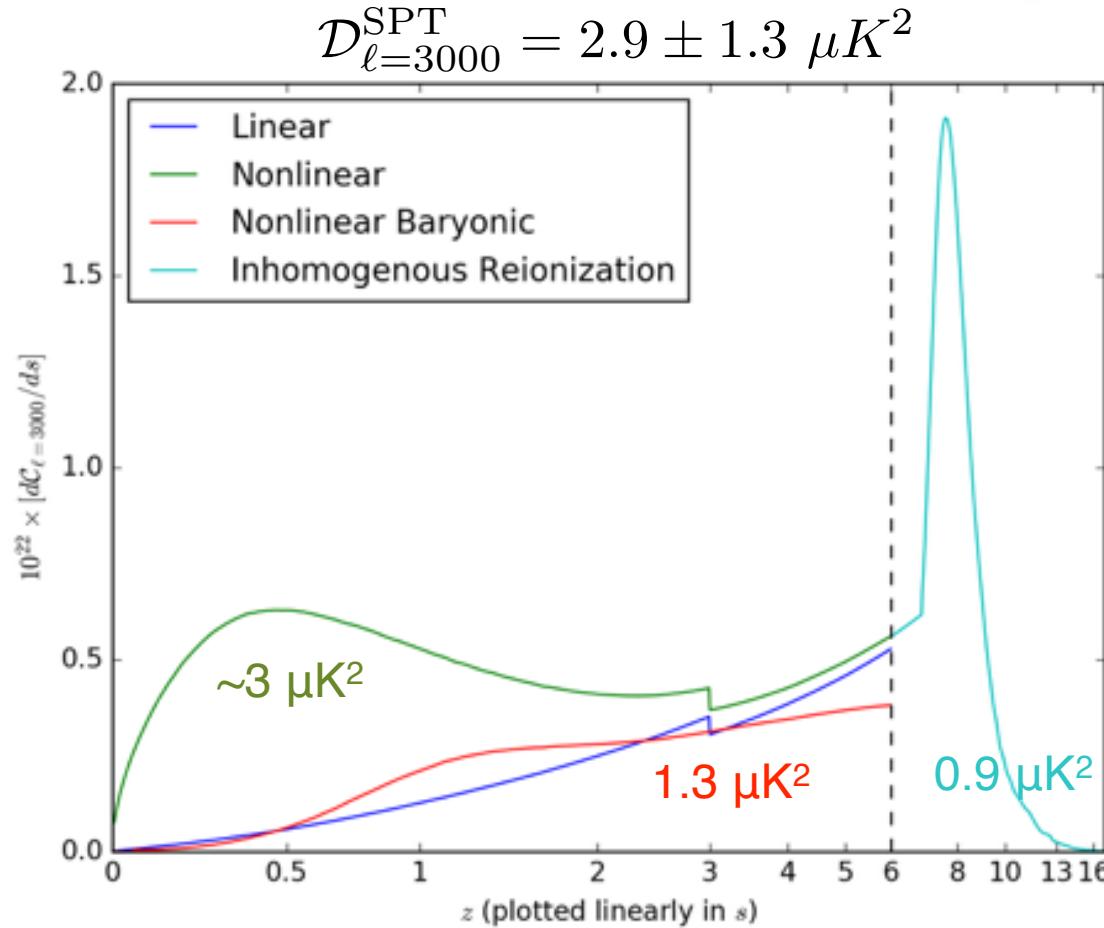


**Extra suppression at small scales and low-redshift due to pressure.**

(Park et al. in prep.)

# Result for Part 3

## Post-reionization kSZ signal



We find  $1.3 \mu\text{K}^2$  for the post-reionization signal accounting for both the effects of nonlinear growth of structures and baryonic physics.

(Park et al. in prep.)

# Summary

- The impact of low-mass galaxies are not distinguishable in the kSZ signal at  $l = 3000$  where we currently observe, but it is at  $l = 10000$ .
- There is a missing term in the expression for  $P_{q,\text{perp}}$  from previous literature. That term adds  $\sim 10\%$  extra to the post-reionization signal.
- Using Illustris simulation we model baryonic effects in the post-reionization ( $z < 6$ ) kSZ signal and include it in the calculation. The result is  $D_{l=3000} = 1.3 \mu\text{K}^2$ .